

Hubble Space Telescope Observations of Nine High-Redshift ESSENCE Supernovae^{1,2,3}

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ABSTRACT

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We present broad-band light curves of nine supernovae ranging in redshift from 0.5 to 0.8. The supernovae were discovered as part of the ESSENCE project, and the light curves are a combination of Cerro Tololo 4-m and *Hubble Space Telescope (HST)* photometry. On the basis of spectra and/or light-curve fitting, eight of these objects are definitely Type Ia supernovae, while the classification of one is problematic. The ESSENCE project is a five-year endeavor to discover about 200 high-redshift Type Ia supernovae, with the goal of tightly constraining the time average of the equation-of-state parameter [$w = p/(\rho c^2)$] of the “dark energy.” To help minimize our systematic errors, all of our ground-based photometry is obtained with the same telescope and instrument. In 2003 the highest-redshift subset of ESSENCE supernovae was selected for detailed study with *HST*. Here we present the first photometric results of the survey. We find that all but one of the ESSENCE SNe have slowly declining light curves, and the sample is not representative of the low-redshift set of ESSENCE Type Ia supernovae. This is unlikely to be a sign of evolution in the population. We attribute the decline-rate distribution of *HST* events to a selection bias at the high-redshift edge of our sample and find that such a bias will infect other magnitude-limited SN Ia searches unless appropriate precautions are taken.

Subject headings: galaxies: distances and redshifts — cosmology: distance scale — supernovae: general

1. Introduction

Type Ia supernovae (SNe Ia) are outstanding probes for cosmology (see Filippenko 2004, 2005, for extensive reviews). Hamuy et al. (1996b) and Riess, Press, & Kirshner (1996) demonstrated that precise values of the Hubble constant can be obtained using SNe Ia as distance calibrators, and that Hubble’s law is linear to a high degree of accuracy at small redshifts. Garnavich et al. (1998a) showed, also using a small sample of high-redshift SNe Ia, that the matter density of the Universe must be considerably less than the critical value in an Einstein-de Sitter universe. Riess et al. (1998) and Perlmutter et al. (1999) found, surprisingly, that SNe Ia at high redshift are systematically “too distant” (by ~ 0.25 mag in distance modulus at a redshift of 0.5 compared to a model of the Universe with $\Omega_M = 0.3$, $\Omega_\Lambda = 0.0$), implying that the Universe is expanding at a progressively greater rate.

This deduction is straightforward to understand. Let us first consider the general form of the “effective” distance¹⁹ of a galaxy (in megaparsecs) (e.g., Longair 1984, Eq. 15.49):

$$d_{\text{eff}} = \frac{2c}{H_0 \Omega_M^2 (1+z)} \left\{ (\Omega_M - 2) \left[(1 + \Omega_M z)^{\frac{1}{2}} - 1 \right] + \Omega_M z \right\} , \quad (1)$$

where c is the speed of light in km s^{-1} , H_0 is the Hubble constant in $\text{km s}^{-1} \text{Mpc}^{-1}$, Ω_M is the mass density of the Universe compared to the critical density, and z is the redshift. The relation assumes a cosmological constant (Λ) of zero. For an empty universe this becomes

$$\lim_{(\Omega_M \rightarrow 0)} d_{\text{eff}} = \frac{cz}{H_0} \frac{(1 + \frac{z}{2})}{(1+z)} . \quad (2)$$

The observed flux (F) of a light source, measured in energy units per unit area per unit time, is related to the luminosity L , effective distance, and redshift as follows:

$$F = \frac{L}{4\pi d_{\text{eff}}^2 (1+z)^2} . \quad (3)$$

The first factor of $(1+z)$ arises because photons produced at frequency ν are observed at frequency $\nu/(1+z)$; that is, they lose energy due to the redshift. We need a second factor of $(1+z)$ because of time dilation of the arrival of the photons.

We may thus define the “luminosity distance” (in Mpc) to be

$$d_{\text{lum}} = d_{\text{eff}}(1+z) . \quad (4)$$

The distance modulus is related to the luminosity distance by the standard equation

$$m - M = 5 \log (d_{\text{lum}} \times 10^6) - 5 = 5 \log (d_{\text{lum}}) + 25 , \quad (5)$$

where the factor of 10^6 is used because cosmological distance is commonly measured in Mpc, not pc.

For SNe, we determine the extinction along their lines of sight and their extinction-corrected rest-frame apparent magnitudes (m) at maximum brightness, and then deduce

¹⁹Carroll, Press, & Turner (1992, Eq. 20) refer to this as the “proper motion distance.” A derivation of Eq. 1 above is given by Krisciunas (1993, Appendix C).

from the light curves the absolute magnitudes (M) at maximum. The calibration of the absolute magnitudes is anchored using nearby SNe Ia whose luminosities and distances are consistent with the value of the Hubble constant used above.²⁰

In units of c/H_0 , an Einstein-de Sitter universe ($\Omega_M = 1, \Omega_\Lambda = 0$) gives luminosity distances which increase as $z + \frac{1}{4}z^2$ over the redshift range of the ESSENCE survey. In the empty-universe model the luminosity distances increase as $z + \frac{1}{2}z^2$. In a universe with $\Omega_M \approx 0, \Omega_\Lambda \approx 1$ the luminosity distances increase approximately as $z + z^2$. Note the progression of the coefficients of z^2 and the order of the consequent luminosity distances. The point is that, depending on the cosmological parameters, the loci fan out in the Hubble diagram, and if the observed distance moduli are larger than one obtains with the empty-universe model, one must consider a positive cosmological constant.

In practice, we plot the differential distance moduli (i.e., observed values minus those for an empty-universe model) vs. the redshift. Distance moduli of SNe up to $z \approx 1.2$ are observed to be systematically larger than we would expect to obtain in an empty-universe model. Riess et al. (2004, Fig. 6) find that SNe Ia have distance modulus differentials which peak at a redshift of 0.46 ± 0.13 . The simplest deduction is that the Universe must contain matter and some form of “dark energy” with a significant negative pressure. The dark energy behaves like a non-zero vacuum energy and causes an acceleration of the expansion.

Grey dust along the line of sight (e.g., Aguirre 1999) or SN Ia evolution in average luminosity over several billion years could explain the implied faintness of SNe at $z \sim 0.5$. Another concern would be selection effects in the discovery of high-redshift SNe. Li, Filippenko, & Riess (2001) and Benetti et al. (2005) discuss the diversity of SNe Ia and how the relative numbers of the different sub-types are affected in magnitude limited surveys. Clocchiatti et al. (2000) and Homeier (2005) discuss the effect of contamination in cosmological surveys by Type Ibc SNe. Examples such as SN 1992ar were even brighter than the brightest SNe Ia at maximum light, but, on average, stripped core SNe are two magnitudes fainter than SNe Ia. Without high quality spectra or three (or more) filter photometry (see, for example, Poznanski et al. 2002; Gal-Yam et al. 2004) it might be difficult to distinguish between Types Ia and Ibc. Since the ESSENCE project relies primarily on two-band photometry and rather “grassy” spectra, our ability to distinguish between SNe Ia and Type Ibc SNe is limited; these limitations apply to other surveys too.

Recently, however, Tonry et al. (2003), Knop et al. (2003), and Barris et al. (2004) have

²⁰In point of fact, we determine the distance *ratio* of a high-redshift SN to the low-redshift limit of the nearby sample. This effectively eliminates the need to know the Hubble constant and the absolute magnitude of a typical SN Ia in the analysis.

tested supernova systematics out to $z \approx 1$, confirming and strengthening the evidence for an acceleration. Riess et al. (2001) and Riess et al. (2004) have pushed the boundary for SN discoveries out to a redshift of $z \approx 1.7$ using the *Hubble Space Telescope (HST)*. At $z \gtrsim 1.2$ the SNe are *brighter* on average than one would expect compared to the empty-universe model. On this basis we can reject the notion of a significant amount of grey dust along the line of sight, as its effect would presumably be even greater at higher redshifts. At $z \gtrsim 1.2$ we observe the Universe when it was small enough that the gravitational attraction of all the matter exceeded the repulsive effective of the dark energy; the Universe was therefore *decelerating*. Indeed, Riess et al. (2004) find that the transition from deceleration to acceleration could have occurred as recently as $z \approx 0.5$.

A key part of the present-day concordance model of the Universe is some form of dark energy. If the equation-of-state parameter $w \equiv p/(\rho c^2) = -1$, where p is the pressure and ρ is the density, then the dark energy is in the form of a standard cosmological constant (e.g., Peebles & Ratra 2003).

Dark-energy models with $w \neq -1$ require internal degrees of freedom, or the presence of non-adiabatic stress perturbations, to remain gravitationally stable (Hu 2005). If $w > -1$, the dark-energy density slowly decreases as the Universe expands; $w \geq -1$ is required by the “weak energy condition” of general relativity (Garnavich et al. 1998b). The case of $w < -1$ was discussed by Carroll (2004). It results in an ever-faster expansion, leading to a “Big Rip.” Such an energy was postulated by Caldwell (2002) and dubbed a “phantom field.”

Upadhye, Ishak, & Steinhardt (2005) present an excellent summary of the current constraints and forecasts for dark energy. They characterize the form of the equation-of-state parameter as $w = w_0 + w_1 z$ out to $z = 1$ and find that existing data for the cosmic microwave background, SNe Ia, and the galaxy power spectrum give $w_0 = -1.38^{+0.30}_{-0.55}$ (2σ), and $w_1 = 1.2^{+0.64}_{-1.06}$ (2σ). After discussing ongoing and planned experiments, they conclude that, “unless we are lucky enough to find a dark energy that is very different from the cosmological constant [$w \equiv -1$], new kinds of measurements or an experiment more sophisticated than those yet conceived will be needed in order to settle the dark-energy issue.”

What is the mass density of the Universe? Eisenstein et al. (2005) investigate the large-scale structure of the Universe using nearly 47,000 luminous red galaxies from the Sloan Digital Sky Survey (SDSS). They find that $\Omega_M = 0.273 \pm 0.025 + 0.123(1+w_0) + 0.137\Omega_K$, and that the curvature term Ω_K is statistically equal to zero (-0.010 ± 0.009). However, Cole et al. (2005) derive a lower value of the mass density; their revised result from over 220,000 galaxy redshifts in the Two-degree Field (2dF) redshift survey is $\Omega_M h = 0.168 \pm 0.016$. With $h \equiv H_0/100 = 0.72 \pm 0.06$ from Freedman et al. (2001), we obtain $\Omega_M = 0.233 \pm 0.030$.

The first results from the Wilkinson Microwave Anisotropy Probe (WMAP) have shown that the spatial geometry of the Universe is flat (Bennett et al. 2003); the total energy density of the Universe corresponds to $\Omega_{tot} = \Omega_M + \Omega_\Lambda = 1.02 \pm 0.02$. Tegmark et al. (2004) use data from WMAP and SDSS to constrain the Hubble constant $H_0 = 70_{-3}^{+4} \text{ km s}^{-1} \text{ Mpc}^{-1}$, the matter component of the Universe $\Omega_M = 0.30 \pm 0.04$, and neutrino masses to less than 0.6 eV. The addition of SNe Ia to the mix gives an age of the Universe of $t_0 = 14.1_{-0.9}^{+1.0} \text{ Gyr}$.

Ever since the luminosities at maximum of SNe Ia were first convincingly shown to be related to the post-maximum rate of decline (Phillips 1993), SNe Ia have come to be regarded as the most common reliable cosmic beacons with which to determine extragalactic distances beyond 20 Mpc. Hamuy et al. (1996c), Phillips et al. (1999), Germany et al. (2004), and Prieto, Rest, & Suntzeff (2005) have refined the $\Delta m_{15}(B)$ method, where the decline-rate parameter is defined to be the number of B -band magnitudes that a SN Ia declines in the first 15 days after maximum. This method previously used the BVI light curves and now uses R -band photometry as well. The “stretch method” (Perlmutter et al. 1997; Goldhaber et al. 2001) scales B -band and V -band templates in the time domain to fit actual light curves. The Multi-color Light-Curve Shape (MLCS) method of Riess, Press, & Kirshner (1996) and Riess et al. (1998) uses the $BVRI$ light curves to give a parameter Δ , the number of magnitudes that a SN Ia is brighter than ($\Delta < 0$) or fainter than ($\Delta > 0$) a fiducial light curve. Jha (2002) and Jha, Riess, & Kirshner (2005) expand the MLCS method to include U -band data. This is very important for studies of high-redshift SNe, because photometry of objects at $z > 0.8$ observed in the R band corresponds to rest-frame photometry in the U band. Finally, Wang et al. (2003) have shown that useful information such as reddening can be derived from plots of the filter-by-filter magnitudes vs. the photometric *colors* (instead of vs. time), using data from the month after maximum light.

The ESSENCE project has as its prime objective the measurement of the time average of the equation-of-state parameter w to an accuracy of $\pm 10\%$. A description of the strategy and methodology of the project is given by Miknaitis et al. (2005). Briefly stated, over a five-year period we shall discover roughly 200 SNe Ia at $z = 0.2$ – 0.8 using the facility CCD mosaic camera on the 4-m Blanco telescope at Cerro Tololo Inter-American Observatory (CTIO). In the past we have observed individual SNe with up to seven telescope/camera combinations. By obtaining all of the ground-based ESSENCE photometry with a single telescope and camera, we should be able to better control systematic photometric errors.

In this paper we present a small fraction of our eventual sample of 200 SNe Ia. Each object presented here was observed with the CTIO 4-m telescope through the R and I bands, and with HST using the Advanced Camera for Surveys (ACS) and its F625W, F775W, and F850LP filters. The nine objects presented here were members of our highest redshift

subsample. Since one of the goals of the ESSENCE project is to identify and minimize sources of systematic error, observing some of our highest redshift objects with *HST* had obvious practical advantages since *HST* gave the highest S/N photometry.

Spectra of the SNe themselves and/or the host galaxies are being obtained with the two Keck telescopes, the VLT, Gemini North and South, both Magellan telescopes, the MMT, and at the Fred L. Whipple Observatory (Mt. Hopkins). Matheson et al. (2005) describe the spectroscopic aspects of the SN search and discuss results obtained thus far using the spectral cross-correlation program SNID (Tonry et al. 2005, in preparation). Suntzeff et al. (2005) have shown that there are insignificant systematic differences (~ 0.02 mag) between *HST* photometry with WFPC and ground-based photometry. We assume that *HST*/ACS has been tested and calibrated sufficiently well that publicly available photometric zeropoints and filter profiles allow us to combine ACS data and ground-based photometry without problems. For one of our objects we also obtained three orbits of data with the *HST* infrared camera NICMOS.

2. Observations

In the first two seasons of the ESSENCE project we discovered 46 definite SNe Ia, 6 likely SNe Ia, 5 core-collapse SNe, plus a number of candidates that were neither confirmed nor rejected as SNe; see Matheson et al. (2005) for a thorough discussion. In the third year of the ESSENCE project we discovered 45 additional SNe, of which 30 are definitely SNe Ia, 10 are possible SNe Ia, and 5 are core-collapse SNe.

Are these relative numbers sensible? Dahlen et al. (2004) considered rates of SNe found in the Great Observatories Origins Deep Survey (GOODS). These authors found 17 SNe Ia to $z = 1.0$ and 16 core-collapse SNe (Types II and Ibc) to $z = 0.9$. Given that GOODS found SNe Ia to $z = 1.6$, we may consider GOODS close to being volume limited at the lower redshifts just stated. This is what is being found with nearby SNe discovered by the Katzman Automatic Imaging Telescope (Leaman, Li, & Filippenko 2004) – comparable numbers of SNe Ia and core-collapse SNe are found in a volume limited sample. ESSENCE, however, is a *magnitude limited* survey. On average, core collapse SNe are at least 2 mag fainter at maximum compared to SNe Ia. If we discover SNe Ia out to $z \sim 0.8$, we discover core-collapse SNe to about $z \sim 0.45$. ESSENCE should be finding between 10 and 20 times as many SNe Ia compared to core-collapse SNe on the basis of the relative volumes being sampled. Since our experiment has as its goal the discovery of 200 SNe Ia, we even try to select against Type II SNe by imposing a color cut on the candidates flagged as potential objects of interest. Since Type II SNe are very blue prior to maximum light, we ignore

candidates with $R - I < 0$. This increases the chance that a spectrum of a candidate will reveal a SN Ia. Thus, the low percentage of core-collapse SNe found in our survey is a measure of our success.

In this paper we report photometry of nine objects discovered in October, November, and December 2003 which were also observed with *HST*/ACS (using the Wide Field Camera). The goal of the *HST* observations was to observe the highest-redshift end of the ESSENCE sample since photometry becomes difficult with the CTIO 4-m as they fade. In Table 1 we list the nine objects. Their redshifts were obtained from observations with Keck I + LRIS, Gemini + GMOS, VLT + FORS1, or Magellan (Baade + IMACS). Spectroscopic details are given by Matheson et al. (2005).

Figure 1 shows *HST*/ACS images, taken through the F775W filter, of the nine SNe discussed in this paper. Each represents the combination of two integrations per date and from four to six dates per object. The total integration times were 3700 s (e510), 5100 s (SN 2003jo), and 4400 s (the other seven SNe). As one can see, only SN 2003jo, SN 2003ll, and SN 2003li were hosted by galaxies sufficiently large and bright to show significant structure. It is not entirely clear that the galaxy near SN 2003kv is the host of that SN. Several of our objects were effectively “hostless” SNe, meaning that the hosts, if they exist, have surface brightness too low to be detected in *HST* and ground-based images.

The program SNID indicates that SN 2003ku is a SN Ia, but there is ambiguity regarding its true redshift. While the most likely redshift on the basis of the spectrum is 0.79, the next most likely redshift is 0.41; see § 3.3 and Matheson et al. (2005) for more comments.

We observed supernova e510 spectroscopically on 29 November 2003 (UT dates are used throughout this paper). We obtained one 1800 s spectrum with Gemini North + GMOS, but were then shut down by high humidity, and we were unable to obtain a redshift from the nearly featureless spectrum. We also attempted to get spectra with Keck, but were hampered by high humidity and clouds. In mid-October 2004, long after the SN had faded, we attempted to get a redshift from spectra of the faint host galaxy. We took three 1800 s spectra with Gemini North + GMOS, but there was insufficient signal to extract a usable spectrum which would provide a redshift. However, following the guidelines of Riess et al. (2001) and Barris & Tonry (2004), one can derive a redshift-independent distance to a *set* of SNe Ia, provided one has sufficiently good photometric coverage in more than a single passband. For any individual object the results are understandably somewhat uncertain. Analysis using the code of Prieto, Rest, & Suntzeff (2005) gives a minimum reduced χ^2 value of the light-curve fits at a redshift of 0.68 for e510. However, any redshift between 0.64 and 0.84 will give a reduced χ^2 value less than 1.0 for this object. It is best if we eliminate it from further consideration. Because there is no spectral information for e510, it never

received an IAU designation.

Matheson et al. (2005) indicate that SN 2003kv was probably a SN Ia, but the spectroscopic identification was not certain. Photometric analysis (below) is entirely consistent with the notion that SN 2003kv was a SN Ia.

The ESSENCE supernovae were not observed as “targets of opportunity” (TOOs) with *HST* since these are very disruptive to the scheduling of the telescope. Instead, a “pseudo-TOO” method was employed which was pioneered by the high-redshift supernova search teams several years ago. We specify months before observation (in the Phase II submission) the region on the sky to be searched, and we specify dates that exact supernova positions will be available for insertion into the *HST* schedule. Usually it is five days to a week before *HST* begins observing the supernovae after the coordinates have been delivered.

Generally we have a number of supernovae in each field from which to choose for *HST* observations, but there are times when a non-optimal target must be inserted at the deadline. Also, because of the delays between the supernova discovery, spectral confirmation, and the *HST* schedule upload, the supernovae are rarely observed before maximum light with *HST*. Since *HST* was meant to improve the quality of photometry at late times for our faintest targets, this delay is not a major problem for this study.

Our *HST/ACS* data were obtained using the F625W, F775W, and F850LP filters, which are essentially the same as the r' , i' , and z' filters of the SDSS photometric system (Smith et al. 2002). The throughput transmission curves with *HST/ACS* are shown in the Appendix. One of our objects was observed in the near-infrared with NICMOS and its F110W filter (similar to a J -band filter).

HST/ACS photometric calibration is based on the zeropoint values of Sirianni et al. (2005), which in turn are based on the Vega spectrophotometric calibration of Bohlin & Gilliland (2004). For an aperture of 50-pixel radius, the Vegamag zeropoints are $z_{pF625W} = 25.731$, $z_{pF775W} = 25.256$, and $z_{pF850LP} = 24.326$ mag (Sirianni et al. 2005, Table 11).

Our *HST/ACS* magnitudes are based on small-aperture photometry (4-pixel radius), using aperture corrections to $r = 50$ pixels consistent with the radial profiles delineated in Table 3 of Sirianni et al. (2005). This aperture photometry was done using the APPHOT package within IRAF.²¹

In the case of SN 2003jo we obtained *HST* template images on 22 May 2004, some 212

²¹IRAF is distributed by the National Optical Astronomy Observatory, which is operated by AURA, Inc. under cooperative agreement with the National Science Foundation.

observer-frame days after the date of maximum light (or 139 rest-frame days). SN 2003jo was faintly visible in the F625W and F775W images on that date, but not visible in the F850LP images. To remove the light of the host galaxy we subtracted the templates, then made corrections (0.02 to 0.06 mag) to the photometry based on the late-time detections.

Template images of SN 2003lh were obtained on 8 June 2004, some 117 rest-frame days after $t(B_{max})$. The SN is visible in the F775W template and undoubtedly is present in the other two. For *HST* photometry of this object, we could perform image subtraction to eliminate any effect of host-galaxy light, then correct for the presence of the SN in the templates using light curves of a different slow decliner such as SN 1991T (Schmidt et al. 1994; Lira et al. 1998). We determined that these corrections would be as large as ~ 0.2 mag. Or, we could simply perform small-aperture photometry (4-pixel radius) without image subtraction. We found that the two methods gave the same final photometric values within 1σ , and we adopt the photometry from the latter method, as it was simpler and relied on fewer assumptions.

HST template images of SN 2003ll were obtained on 5 October 2004. The other objects discussed in this paper showed no significant host-galaxy light near the locations of the SNe in the *HST* images.

In Table 2 we list the photometry of the SNe using ACS/WFC with *HST*. We also give three orbits of F110W data for one SN, obtained with *HST* and its near-infrared camera NICMOS.

The ground-based images with the CTIO 4-m telescope were taken through *R*-band and *I*-band filters, the details of which are described in the Appendix. Our *R* filter is essentially the same as the Bessell (1990) *R* filter, but our *I*-band filter has steeper short-wavelength and long-wavelength cutoffs than Bessell’s *I*-band filter.

We found that the ground-based imaging in the *R* band gave higher signal-to-noise ratio (S/N) detections for the objects with $z \lesssim 0.6$, but that the *I*-band imaging gave higher S/N detections for the objects with $z \gtrsim 0.6$. This is just an empirical consequence of the spectral energy distributions of SNe Ia in the rest-frame *UBV* bands coupled with the quantum efficiency of the CTIO 4-m mosaic camera in *R* and *I*.

Ground-based photometry was carried out using template images obtained at least 18 *rest-frame* days prior to the observed maximum, or, in one case, long after the SN had faded. (At 18 to 20 rest-frame days prior to maximum light, any light of a high- z SN on the rise would be lost in the sky noise of the CTIO 4-m images.) When possible we used a median of three template images obtained on photometric nights to eliminate cosmic rays in the templates. For eight of the nine objects discussed here we used reference images from early

in the 2003 observing season. In the case of SN 2003jo we used reference images from 28 November 2002 (R band) and 11 October 2004 (I band).

For the rotation, alignment, kernel matching, and difference imaging of the ground-based images we used two packages of scripts written by one of us (BPS). The result is point-spread function (PSF) magnitudes of field stars and the SNe themselves. The scripts rely on the kernel-matching algorithm of Alard & Lupton (1998). For reasonably high S/N detections of the SNe, we adopted the 1σ symmetrical error bars in magnitudes from DOPHOT. For faint signals, the error bars in magnitude space are not symmetrical. We then chose a medium-sized number (20) of random locations in the subtracted images (avoiding obvious image defects) to derive the sky noise, and derived 1σ error bars in flux space, which we then converted to asymmetrical upper and lower errors in magnitude space. In the case of the CTIO 4-m R -band images, the sky noise is roughly 250 analog-to-digital units (ADUs, or counts, where 1 ADU $\approx 1.8 e^-$) in 200 s exposures (templates and data images) on clear, moonless nights. 400 s I -band images give corresponding sky noise of about 450 ADUs in the subtracted images. A SN Ia at maximum light and $z \approx 0.5$ typically gives a signal of 4000 ADUs in corresponding R -band and I -band exposures obtained with the CTIO 4-m telescope.

We used data from the early data release of the SDSS to obtain R -band and I -band magnitudes for the field stars near the SNe discussed in this paper, relying on the transformations given in Table 7 of Smith et al. (2002). We avoided using field stars with $r' - i' > 0.95$ mag, as many stars this red are variable and the scatter in the photometric transformation becomes large.

To check our algorithm for calibrating the SN photometry, we obtained images of six of our nine SN fields on 20 October 2004 using the CTIO 0.9-m telescope. This was a photometric night, allowing the determination of the atmospheric extinction values and the instrumental color terms using observations of seven Landolt (1992) fields. From aperture photometry with an 8 pixel (3.2'') radius aperture, we obtained sensible photometry to $R = 19.7$ mag with 600 s to 720 s exposures. In 900 s I -band exposures we effectively reached magnitude 19.0. This provided ~ 1.6 mag of overlap with the brightest unsaturated stars in our CTIO 4-m images (which were typically 200 s in R and 400 s in I). A comparison of the derived RI magnitudes and the values obtained from the transformation from SDSS $g'r'i'$ magnitudes is shown in Figure 2. The R -band differentials give a slope of -0.0049 ± 0.0031 mag per mag, while the I -band differentials give a slope of -0.0087 ± 0.0058 . Δ magnitude is in the sense “observed values from CTIO 0.9-m photometry directly tied to Landolt (1992) standards” *minus* “values derived from SDSS photometry”. Neither slope differs from zero at a statistically significant level. At $R = 19.0$ mag, $\langle \Delta R \rangle = -0.027 \pm 0.051$ mag, while at

$I = 19.0$ mag, $\langle \Delta I \rangle = -0.025 \pm 0.046$ mag. Neither differs significantly from zero. We are trying to account for all sources of systematic error greater than 0.01 mag in the ESSENCE photometry. Hence, a more robust test involving more fields and nights is warranted.

In Table 3 we give the photometry of the SNe using the CTIO 4-m telescope and its facility mosaic camera. We list the R and I magnitudes in the natural magnitude system of the telescope/filter/detector transmission function, with a zero point based on the Landolt (1992) system. To force the zero point of the natural system to that of Landolt, we plot $(R_{\text{nat}} - R_{\text{Landolt}})$ as a function of $(V - R)$ and $(I_{\text{nat}} - I_{\text{Landolt}})$ as a function of $(V - I)$. By forcing $(R_{\text{nat}} - R_{\text{Landolt}})$ and $(I_{\text{nat}} - I_{\text{Landolt}})$ to be 0.00 where $(V - R)$ and $(V - I)$ are zero, we transfer the zero point of the Landolt system onto the natural system. See the Appendix of Schmidt et al. (1998).

In Figures 3 and 4 we show the observed light curves of the nine SNe discussed here. Because the central wavelengths of the R -band and F625W filters are similar, we would expect that photometry in those bands would be reasonably similar. Also, the I -band photometry should be similar to the F775W and F850LP photometry. An exception occurs with R -band and F625W photometry if we are observing a SN Ia with $z \approx 0.8$; the flux of the SN just longward of the Ca II H & K lines is included in the R -band photometry but excluded from the F625W band. This can make a difference of ~ 0.3 mag three weeks after maximum light in the observer's frame.

3. Discussion

3.1. Light-Curve Fits and a Composite Spectrum

Seven of the nine objects discussed in this paper have unambiguous redshifts on the basis of spectra (Matheson et al. 2005). We were unable to obtain a redshift of e510 from the SN itself when it was visible, or from the faint host galaxy the following year, so it was never assigned an official name by the IAU. But its light-curve shapes and maximum magnitudes are completely compatible with those of other high-redshift SNe Ia. The case of SN 2003ku is discussed below.

In order to derive the distances of the SNe, we first K-corrected the ground-based and space-based photometry to rest-frame U , B , V , or R photometric bands. For the MLCS method of Jha, Riess, & Kirshner (2005, hereafter MLCS2k2) and the $\Delta m_{15}(B)$ analysis (Prieto, Rest, & Suntzeff 2005), if $z \leq 0.6$ the ground-based R -band and F625W data were K-corrected to rest-frame B , ground-based I -band and F775W data were transformed to rest-frame V , and F850LP data were transformed to rest-frame R . Photometry of SN 2003kp was

transformed to B and V . For SN 2003kv we transformed the photometry to UBV instead of BVR .

In Tables 4, 5, and 6, we give the light-curve fits using MLCS2k2, the Bayesian Adapted Template Method (BATM; Tonry et al. 2003, and Tonry et al. 2005, in preparation), and the $\Delta m_{15}(B)$ method of Prieto, Rest, & Suntzeff (2005), respectively. Our derived distance moduli are consistent within the errors using the three light-curve fitting methods. In all three tables we give the differences of the derived distance moduli and the corresponding values in an empty universe ($\Omega_M = 0.0, \Omega_\Lambda = 0.0$).

Since the light curve decline of SN 2003jo was covered by the ground based data, we also fit this object without the HST data. Using just the CTIO RI data, the method of Prieto, Rest, & Suntzeff (2005) gives $\Delta m_{15}(B) = 0.84 \pm 0.17$, $E(B - V) = 0.05 \pm 0.05$, and $(m - M) = 42.77 \pm 0.20$ (on an $H_0 = 65$ scale). These values are consistent with the solution given in Table 6, which used the combined CTIO 4-m and HST data.

The $\Delta m_{15}(B)$ values given in Table 6 indicate that all but one of the objects discussed here are slow decliners. The slowest declining template object in the Prieto et al. training set is SN 1999aa, with $\Delta m_{15}(B) = 0.81 \pm 0.04$. Thus, our objects are near, but not beyond, the limit of the $\Delta m_{15}(B)$ system.

The MLCS fits were done in two ways. First, we assumed a prior that no light curve could be slower than MLCS $\Delta = -0.4$, the minimum value of the MLCS training set. In this case we found that the data systematically deviated from the fits at late times. The actual light curves were slower than the slowest declining objects in the training set. We then fit the light curves with no constraint on the possible values of Δ . In this case we had to extrapolate beyond the training set. In §3.4 we discuss the cosmological effects of using the prior or not using it.

In Figure 5 we show the $\Delta m_{15}(B)$ light-curve fits in the rest-frame bands for 7 SNe. In Figure 6 we also show the K-corrected, extinction-corrected I -band data of SN 2003lh, our only data obtained with NICMOS. For comparison we show the I -band data of SN 1999aa (Krisciunas et al. 2000; Jha 2002), the slowest decliner in the nearby sample used by Prieto, Rest, & Suntzeff (2005), and the I -band data of the prototypical slow decliner SN 1991T (Lira et al. 1998). The photometry of SN 1999aa has been adjusted in magnitude space to the brightness of SN 2003lh using the $\Delta m_{15}(B)$ distance modulus given in Table 6 and assuming an absolute magnitude $M_I(\text{max}) = -19.1$ for the slowest-declining SN Ia studied by Nobili et al. (2005). For $H_0 = 65 \text{ km s}^{-1} \text{ Mpc}^{-1}$, the value used for the $\Delta m_{15}(B)$ system, this becomes $M_I = -19.32$ mag. The maximum of SN 1991T is made to coincide with that of SN 1999aa. At face value the I -band secondary hump of SN 2003lh was weaker than that

of SNe 1991T and 1999aa. As we pointed out in a previous paper (Krisciunas et al. 2001, Fig.18), SNe Ia with identical decline rates in B and V can have significantly different I -band secondary maxima. Whatever is the appropriate adjustment of the SN 2003lh photometry in Figure 6, we can say that its secondary I -band hump occurred earlier than that of SN 1999aa.

As a measure of our systematics, we show in Figure 7 the residuals of the K-corrected CTIO 4-m and *HST* data compared to the rest-frame MLCS2k2 light-curve fits. In Figure 8 we show analogous plots of the residuals of the $\Delta m_{15}(B)$ fits shown in Figure 5. In these two figures a differential magnitude greater than zero means that a rest-frame datum is fainter than the fit, and a differential magnitude less than zero means that a rest-frame datum is brighter than the fit.

Figures 7 and 8 show that there are no statistically significant trends in the residuals of the K-corrected data compared to the light-curve fits, and that the mean residual is close to zero, as expected.

For the individual spectra of the SNe Ia discussed in this paper, see Matheson et al. (2005). In an attempt to find spectral peculiarities of the slow decliners consistent with their light curves, we created a composite spectrum of SNe 2003jo, 2003kp, 2003kv, 2003lh, 2003le, and 2003li (see Figure 9). We first deredshifted each spectrum to its rest frame, then averaged each wavelength bin, using only the spectra which covered that particular wavelength bin. Using the dates of B -band maximum in Table 4, the composite spectrum has a mean spectral age of +5 days with respect to maximum light.

The composite spectrum shows features typical of a normal SN Ia slightly past $T(B_{max})$. The Ca II H&K lines are present and as strong as in the normal SN Ia 1992A at a similar age, unlike the over-luminous, slowly declining SN 1991T (Filippenko et al. 1992). The spectrum of the slowly declining SN 1999aa evolved similarly to SN 1991T, with the major exception of having strong Ca II H&K. At a similar age to the composite spectrum, SN 1999aa does not appear to be drastically different than the composite spectrum.

Despite the increase of S/N in the ESSENCE objects by combining them into one composite spectrum, the overall S/N is lower than the S/N obtained for the spectra of many low- z SNe. We note, however, that the $\lambda 4130$ feature due to Si II is weak and more like SN 1999aa than SN 1992A, consistent with the slowly declining light curve fits found for these SNe. In a future paper we shall provide a full length discussion of composite ESSENCE spectra (Foley et al. 2006).

3.2. Decline Rates of Supernovae

In Figure 10 we show a histogram of $\Delta m_{15}(B)$ values for 107 nearby SNe Ia, dividing them into two groups according to the host-galaxy type (Gallagher et al. 2005). Of these 107 hosts, 71 are spirals and 36 are ellipticals or S0 galaxies. We also add to the histogram seven values for ESSENCE SNe listed in Table 6. Figure 10 shows that spiral galaxies are more likely to produce slowly declining SNe Ia, while ellipticals are more likely to produce rapidly declining ones; for further details see Hamuy et al. (1996a) and Gallagher et al. (2005). Umeda et al. (1999) discuss why this might be the case, but for now it must just stand as an empirical fact. The striking aspect of the histogram is that all but one of the ESSENCE SNe are at the extreme slow-declining end of the $\Delta m_{15}(B)$ distribution.

Figure 1 suggests that SN 2003jo, SN 2003ll, and SN 2003li occurred in spiral galaxies on the basis of their morphology in the *HST* images. SN 2003kv is projected near a bright galaxy, but spectroscopy is needed to show that it is the host. The other five SNe discussed here occurred in very faint hosts, about which we have almost no information for the purposes of morphological or spectroscopic classification. None of the 9 SNe was found in a galaxy that is obviously an elliptical.

Using the local sample as a guide, we would expect a small number (3) of our ESSENCE sample to be in E/S0 galaxies. But such an extrapolation to a magnitude limited sample is dangerous. SNe Ia in nearby early-type galaxies are, on average, fast decliners and intrinsically faint at maximum, so less likely to be found in our search. Therefore, a lack of early-type hosts is not surprising. A more detailed analysis of this bias is discussed below.

For a discussion of the properties of nearby galaxies that have hosted SNe Ia, see Gallagher et al. (2005). The morphology of the host galaxies of high- z SNe Ia has been discussed by a number of authors. Farrah et al. (2002) studied 22 galaxies at $z \approx 0.6$ and found that $\sim 70\%$ of high- z SNe occur in spirals and $\sim 30\%$ occur in ellipticals,²² similar to the percentages for the local sample. They found no evidence that SNe Ia are preferentially found in the outer regions of the hosts, implying that host galaxy extinctions of the high- z sample should be comparable to the local examples. Williams et al. (2003) found no correlations of the distance residuals with host-galaxy properties in the redshift range $0.42 < z < 1.06$. Their 18 galaxies which hosted SNe were discovered by the High- z Supernova Search Team (Schmidt et al. 1998; Riess et al. 1998). Sullivan et al. (2003) found from a larger sample (Perlmutter et al. 1999, 42 objects discovered by the Supernova Cosmology Project) that

²²At least one galaxy in the Farrah et al. sample (the host of SN 1998M) is mis-classified as an elliptical. Our multi-band imagery indicates that the host is a blue star-forming galaxy.

there is evidence for a larger *scatter* of the absolute magnitudes for the high- z SNe occurring in spirals. They also found that SNe occurring in spirals are on average marginally less luminous than those in E/S0 galaxies, by 0.14 ± 0.09 mag. But this is opposite of what is seen at low redshifts (Hamuy et al. 1996a).

These analyses do not lead us to believe that there are significant differences between the hosts of low- z and our high- z SNe Ia or between the SNe Ia themselves. Why, then, do all but one of the ESSENCE objects discussed in this paper have such slow decline rates? Some of the issues involved have already been discussed by Li, Filippenko, & Riess (2001). It is important to note that the SNe Ia in this paper not typical of ESSENCE SNe in general. They were chosen to be at the highest redshift end of the ESSENCE distribution for *HST* follow-up.

In Figure 11 we provide an illustrative answer. Let us assume that we discover all the SNe in the R band, and ignore the I band. Assume a fixed magnitude limit for detection; we shall use $R = 23.0$ mag, which gives $S/N \approx 10$ for images obtained in good seeing ($0.9''$). Varying the limit will change the details but not the general result. We calculate the time a SN Ia stays above the detection limit as a function of $\Delta m_{15}(B)$ and redshift. This is the “control time,” but we will only include the time before maximum, since detection before the time of maximum was a requirement for *HST* follow-up observations. Figure 11 shows that the control time is a steeply falling function of $\Delta m_{15}(B)$. For $z < 0.5$ this is not a problem for the ESSENCE search since we visit the same piece of sky every four nights. But beyond $z \approx 0.6$ the fast-declining events are not above our detection threshold long enough to have been included in the *HST* sample. As we push to high redshift it is clear that our *HST* selection is highly biased to the slowest-declining SNe Ia.

The selection effect which increased our sample of slowly declining light curves relative to the average distribution will affect their cosmological use. Our SNe Ia may be slower than average for their luminosity or intrinsically brighter than expected for their observed light curve shape as a result of the selection effect. While this selection bias is strong in our *HST* sample, it will be present in the highest-redshift slice of any magnitude-limited search for SNe Ia. Fast-declining events are not only fainter than slowly declining SNe Ia, but are above any threshold for less time, leading to a control-time bias against their discovery. In designing a search that minimizes systematic errors, care must be taken to avoid this bias by either visiting fields at a rapid cadence or pruning the highest-redshift end of the accumulated sample. Such a selection effect should be modeled once we have our full sample of 200 SNe Ia to fully remove its impact.

3.3. The Strange Case of SN 2003ku

The spectral cross-correlation program SNID indicated that the spectrum of SN 2003ku was most consistent with that of a SN Ia with $z = 0.79$, but redshift 0.41 was almost as likely (Matheson et al. 2005). In Figure 12 we show these two possibilities. SN 2003ku was the brightest of the nine SNe discussed in this paper (see Figs. 3 and 4); thus, photometric considerations alone would favor the smaller redshift. Using the method of Prieto, Rest, & Suntzeff (2005), our attempts to K-correct and fit the photometry of SN 2003ku to rest-frame B and V magnitudes gave $\chi^2_\nu = 2.8$ for the lower redshift and $\chi^2_\nu = 5.2$ for the larger value, implying that the lower redshift is favored, but the fit is still not very good. If $z = 0.79$, this object gives a derived distance modulus 2.5 mag “too close” compared to the empty-universe model. The BATM analysis gave similar results; if $z = 0.79$, the derived distance modulus is 1.7 mag smaller than one would get in an empty-universe model.

While SNe exhibit some dispersion in luminosities, and light-curve fitting of low S/N photometry adds to the scatter of measurable parameters, something is clearly amiss with an object that is discrepant by 2 mag. For SN 2003ku several possibilities can be considered:

(1) The photometry contains some serious calibration error. One of the fields represented in Figure 2 was the SN 2003ku field. There is a reasonable match, certainly better than 2 mag, of the RI photometry and the HST photometry. We believe the photometry of this object is correct. We note that SN 2003ku was the only object discussed here whose R and F625W photometry *clearly* differ, by ~ 0.4 mag. See Figure 3. From a cursory inspection of Figures 12, 17, and 18 one can see that the local peak in the spectrum at 7100–7300 Å would be included in R -band photometry but excluded from F625W photometry.

(2) Could SN 2003ku be a gravitationally lensed object? Holz (2001) indicates that 0.05% of sources at $z = 1$ are expected to be multiply imaged on arcsecond scales. Porciani & Madau (2000) investigated the lensing magnification of high-redshift SNe for a variety of scenarios and found that up to 10% could be magnified by 0.1 to 0.3 mag at $z = 1$. Further considerations are elaborated by Wang (2005, and references therein). At a redshift of 0.4 to 0.8 we conclude that the probability of a 2 mag magnification is extremely small. Also, there is no evidence of multiple images.

(3) Was SN 2003ku a spectroscopically peculiar SN Ia? This would also help explain the confusion on the part of SNID to determine the redshift.

(4) If it were a SN Ia at $z \approx 0.4$ or less, then it would not have been so overluminous.

(5) It could have been a SN Ia with a different explosion mechanism. Wilson & Mathews (2004), for example, describe scenarios whereby white-dwarf stars passing close to black holes

of a range of masses (10 to $10^9 M_\odot$) explode.²³ These authors estimate that between 10^{-2} and 10^{-4} of the SNe Ia out to $z = 1$ could be white dwarfs disrupted by black holes. Most significantly, they suggest that the light curves would have different shapes. We wonder if white dwarfs disrupted by black holes might also have different spectra, or luminosities which differ from the standard white-dwarf-plus-donor-star scenario.

(6) It could have been an extremely luminous SN Ib or SN Ic like SN 1992ar (Clocchiatti et al. 2000), and with a spectrum different from any in SNID’s database of spectra. We feel this is the mostly likely possibility.

In any case, given the ambiguity of the redshift (hence luminosity) and uncertainty in its classification, SN 2003ku will be eliminated from our final determination of the equation-of-state parameter.

3.4. Cosmological Consequences

We consider three cosmological models: the empty universe ($\Omega_M = 0.0, \Omega_\Lambda = 0.0$), the open universe ($\Omega_M = 0.3, \Omega_\Lambda = 0.0$), and the concordance model ($\Omega_M = 0.3, \Omega_\Lambda = 0.7$). In Figure 13 we show a differential Hubble diagram derived from the light-curve fits, and compare the results to these three models. For each SN shown in the plot we determined the distance modulus using MLCS2k2 and subtracted off the distance modulus one would get in an empty universe. We used the 157 SNe in the “gold” set of Riess et al. (2004) and also included seven ESSENCE objects discussed here. Because the redshifts of SN 2003ku and e510 are uncertain or unknown, we do not include them in the plot. Consider Figure 13. In an empty universe the data points would equally likely fall above and below the horizontal line in the middle, but it is obvious that the majority of the points at $z \approx 0.5$ and $z \approx 0.8$ are above this. The dashed line corresponds to the concordance model, while the dotted line corresponds to the open universe with $\Omega_M = 0.3$. There is a half-magnitude range (a root-mean-square uncertainty of ± 0.18 mag) in the nearby sample ($z < 0.10$). The more distant objects do not show a *considerably* greater range, which is reassuring, since we are assuming that SNe Ia with lookback times of several billion years are similar to those observed nearby.

In Figure 13 the weighted mean difference of the ESSENCE data points, compared to the open-universe model (i.e., dotted line), is $\langle \Delta(m - M) \rangle = +0.37 \pm 0.09$ mag. At $z = 0.6$

²³Khokhlov, Novikov, & Pethick (1993) and Diener et al. (1997) previously described the interactions and disruption of stars, in particular $n = 1.5$ polytropes, with black holes.

the difference between the concordance model (dashed line) and the open-universe model is $+0.23$ mag. In a flat universe, as Ω_M becomes smaller and Ω_Λ increases, the curve bows upward at $z = 0.6$. Thus, if the geometry of the Universe is flat, the ESSENCE data alone would stipulate that the mass density of the Universe is $\Omega_M < 0.3$ and the dark energy has $\Omega_\Lambda > 0.7$.

Using the MLCS2k2 fits and a nearby sample to establish the zeropoint of the distance moduli, we obtain cosmological constraints for a cosmological constant based on the ESSENCE supernovae. In Figure 14 we have plotted constraints assuming a prior limiting the decline rate of the slowest supernovae, and constraints based on no prior which allow an extrapolation by MLCS. Clearly the choice of a prior affects the cosmological results and points to the need to avoid a selection bias that preferentially discovers slowly declining supernovae at the limits of the decline rate range of the local sample. Using the 2dF mass constraint of Cole et al. (2005) ($\Omega_M h = 0.168 \pm 0.016$), the Hubble constant of Freedman et al. (2001), and with the prior on MLCS Δ we find $\Omega_\Lambda = 1.26 \pm 0.18$ for the sample of ESSENCE plus nearby SNe. Without the prior we find $\Omega_\Lambda = 0.99 \pm 0.21$.

In Figures 15 and 16 we show constraints obtained for Ω_Λ and w using the entire “gold” set of 157 SNe from Riess et al. (2004) plus seven ESSENCE SNe discussed in this paper. We have assumed $w = -1$ for Figure 15 and have used two different matter constraints: the 2dF mass constraint of Cole et al. (2005) coupled with the Hubble constant from Freedman et al. (2001), and the SDSS large-scale structure result of $\Omega_M = 0.273 \pm 0.025 + 0.137\Omega_K$ (Eisenstein et al. 2005). Figure 15 shows that the standard model is recovered at the 1σ level if we use the full “gold” SN sample, the seven ESSENCE objects from Table 4, and either matter constraint. From Figure 16 we find $w = -0.88 \pm 0.11$ with the prior on MLCS Δ for the ESSENCE sample. Without the prior we find $w = -0.90 \pm 0.12$. This is a significant shift in w considering only 7 of 164 SNe Ia were affected by the change of prior. This demonstrates the importance of the selection bias at the high- z end of any survey.

4. Conclusions

We have presented photometry of nine supernovae from the ESSENCE project. Ground-based photometry allowed us to cover the maxima in the rest-frame B and V light curves of all the objects discussed here, while the photometry obtained with HST/ACS allowed us to characterize the light-curve tails. The light-curve fitting of seven objects with reliable redshifts, carried out with three different methods, gave distance moduli consistent within the errors.

On the basis of the values of $\Delta m_{15}(B)$ derived for the ESSENCE SNe, all but one are slow decliners compared to the local sample and their light curves are as slow as the slowest found in the local set of well-observed supernovae. We show that the SNe selected to be observed with *HST* were at the high- z end of our distribution of redshifts and probably represent a selection bias for slow-declining events.

The ESSENCE project is being carried out every other night on the CTIO 4-m telescope during the months of October, November, and December. After three years of the ESSENCE project we have discovered roughly 100 SNe and SN candidates. Eventually, with ~ 200 spectroscopically confirmed ESSENCE SNe Ia, all observed with the same ground-based telescope and filter system, we should be able to determine the time average of the equation-of-state parameter of the Universe to $\pm 10\%$.

We observe two sets of fields on alternating observing nights. The resulting cadence of light-curve points every four observer-frame days is clearly sufficient to characterize the light curves. For SNe Ia with $z \approx 0.5$, we obtain apparent magnitudes at maximum with an uncertainty of ± 0.06 mag. For further details on the project strategy see Miknaitis et al. (2005).

The 157 “gold” SNe Ia of Riess et al. (2004), along with seven high-redshift SNe discussed here, give contours in the $\Omega_M - \Omega_\Lambda$ plane consistent with a positive cosmological constant and flat geometry. Further cosmological tests await the acquisition of larger self-consistent data sets.

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Gallagher for providing his database of supernova properties. Lou Strolger kindly derived the F110W photometry of one of our supernovae. K.K. thanks David Spergel for many stimulating discussions. Finally, we thank an anonymous referee for constructive suggestions and references to other work.

5. Appendix: Filters for Supernova Photometry

In Figure 17 we show the effective throughput (i.e., combination of filter transmission and quantum efficiency as a function of wavelength) of the three filters used with *HST*/ACS.

We used laboratory hardware and software produced by Ocean Optics to make transmission-curve traces of the *R*-band and *I*-band filters, which were used with the CTIO 4-m telescope and its facility mosaic camera. The traces were taken with the incidence angle of the input laser beam ranging from 0° to 11° . We subsequently used the data files and a program kindly provided by G. Jacoby to simulate the filter traces appropriate for the $f/2.7$ beam of the CTIO 4-m telescope and its Mosaic camera. As the *I*-band filter is an interference filter, this is particularly important. The effective *I*-band filter profile for an $f/2.7$ beam has half-power points shifted roughly 30 \AA toward the blue from the 0° incidence angle tracing.

The *RI* filter transmission curves can be obtained in graphical and tabular form at website <http://www.ctio.noao.edu/~points/FILTERS>. Our effective *RI* filter transmission curves are shown in Figure 18. They include the effects of reflection off the primary mirror, the quantum efficiency of the CCD chips as a function of wavelength, the atmospheric extinction, and the major telluric absorption lines.

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Table 1. High-Redshift Supernovae: Basic Data

IAU name	ESSENCE	α (J2000)	δ (J2000)	Redshift ^a	$E(B - V)_{Gal}^b$
SN 2003jo	d033	23:25:24.03	−09:26:00.6	0.53	0.036
SN 2003kp	e147	02:31:02.64	−08:39:50.8	0.64	0.032
SN 2003ku	e315	01:08:36.25	−00:33:20.8	0.79 ^c	0.036
...	e510	23:30:59.97	−08:37:34.4	0.68 ^d	0.032
SN 2003kv	e531	02:09:42.52	−03:46:48.6	0.78	0.023
SN 2003lh	f011	02:10:19.51	−04:59:32.3	0.54	0.020
SN 2003le	f041	01:08:08.73	+00:27:09.7	0.56	0.029
SN 2003ll	f216	02:35:41.19	−08:06:29.6	0.60	0.033
SN 2003li	f244	02:27:47.29	−07:33:46.2	0.54	0.027

^aObtained from the spectra of the SNe themselves, rather than the host galaxies. We used our program SNID (Tonry et al. 2005, in preparation).

^bFrom the reddening maps of Schlegel et al. (1998); magnitude units.

^cSee text and Matheson et al. (2005) for a discussion of the redshift of this object.

^dThe minimum reduced χ^2 value of the light-curve fits is obtained for $z = 0.68$; no spectroscopic redshift was obtained.

Table 2. *HST* Photometry^a

SN name	JD ^b	F625W	F775W	F850LP	F110W
2003jo	2946.95	23.031 (0.018)	22.661 (0.017)	22.455 (0.023)	...
...	2953.01	23.399 (0.022)	22.841 (0.018)	22.748 (0.035)	...
...	2960.27	...	23.201 (0.019)
...	2973.81	...	23.799 (0.028)	23.186 (0.100)	...
...	2976.41	25.045 (0.065)	23.863 (0.035)	23.210 (0.035)	...
...	2985.99	...	24.277 (0.040)	23.464 (0.036)	...
...	3148.26	26.75 (+0.51/-0.35)	26.91 (+0.35/-0.26)
2003kp	2981.89	23.540 (0.024)	23.006 (0.020)	22.789 (0.027)	...
...	2988.67	23.961 (0.034)	23.394 (0.025)	23.081 (0.032)	...
...	2995.16	...	23.657 (0.024)	23.282 (0.042)	...
...	3007.67	...	24.198 (0.037)	23.576 (0.038)	...
...	3021.02	...	24.766 (0.053)	23.913 (0.049)	...
2003ku	2976.02	23.145 (0.027)	22.249 (0.015)	21.904 (0.018)	...
...	2983.15	23.614 (0.026)	22.550 (0.017)	22.009 (0.018)	...
...	2989.67	...	22.945 (0.023)	22.114 (0.017)	...
...	3001.88	...	23.601 (0.025)	22.522 (0.021)	...
...	3016.29	...	24.168 (0.053)	23.200 (0.032)	...
e510	2983.08	23.799 (0.043)	23.413 (0.027)	23.315 (0.038)	...
...	2989.75	...	23.723 (0.028)	23.524 (0.026)	...
...	3001.53	...	24.474 (0.051)	23.824 (0.051)	...
...	3016.15	...	25.168 (0.095)	24.298 (0.110)	...
2003kv	2981.82	24.277 (0.042)	23.228 (0.024)	23.149 (0.037)	...
...	2988.81	24.813 (0.065)	23.518 (0.030)	23.379 (0.044)	...
...	2995.09	...	23.896 (0.031)	23.559 (0.040)	...
...	3006.82	...	24.510 (0.049)	24.012 (0.059)	...
...	3020.95	...	25.265 (0.093)	24.592 (0.094)	...
2003lh	3009.76	23.877 (0.031)	23.060 (0.021)	22.979 (0.030)	...
...	3015.96	24.325 (0.060)	23.363 (0.025)	23.119 (0.046)	...
...	3016.60	23.292 (0.026)
...	3022.89	...	23.704 (0.026)	23.179 (0.030)	...
...	3024.36	23.400 (0.028)

Table 2—Continued

SN name	JD ^b	F625W	F775W	F850LP	F110W
...	3030.60	23.463 (0.021)
...	3035.89	...	24.208 (0.037)	23.624 (0.040)	...
...	3042.75	...	24.430 (0.044)	23.888 (0.067)	...
2003le	3003.01	22.899 (0.017)	22.529 (0.017)	22.407 (0.022)	...
...	3009.68	23.299 (0.022)	22.725 (0.018)	22.603 (0.025)	...
...	3016.68	...	22.983 (0.018)	22.814 (0.025)	...
...	3029.68	...	23.501 (0.025)	23.048 (0.030)	...
...	3042.81	...	24.084 (0.041)	23.372 (0.055)	...
2003ll	3004.71	24.593 (0.051)	23.827 (0.033)	23.174 (0.033)	...
...	3011.65	25.351 (0.139)	24.529 (0.055)	23.452 (0.042)	...
...	3018.44	...	24.570 (0.065)	23.496 (0.035)	...
...	3029.75	...	25.070 (0.135)	23.761 (0.040)	...
...	3043.71	...	25.336 (0.083)	24.356 (0.156)	...
2003li	3011.22	23.983 (0.047)	23.121 (0.022)	23.094 (0.034)	...
...	3018.82	24.544 (0.050)	23.459 (0.027)	23.172 (0.035)	...
...	3025.95	...	23.735 (0.037)	23.334 (0.046)	...
...	3036.20	...	24.244 (0.038)	23.577 (0.040)	...
...	3050.60	...	24.705 (0.057)	23.956 (0.053)	...

^aThe ACS magnitudes given are “Vega” magnitudes derived using 50-pixel-radius zeropoints (Sirianni et al. 2005), which are based on the Vega spectrophotometric calibration of Bohlin & Gilliland (2004). The values in parentheses are 1σ error bars. The F110W photometry was obtained with NICMOS.

^bJulian Date minus 2,450,000.

Table 3. Ground-Based R and I Photometry^a

SN name	JD ^b	R_{nat}	I_{nat}
2003jo	2931.52	22.63 (± 0.06)	22.57 ($+0.13/-0.12$)
	2934.51	22.63 (± 0.10)	22.56 ($+0.13/-0.12$)
	2940.54	22.76 (± 0.08)	22.74 ($+0.16/-0.14$)
	2944.51	23.03 ($+0.20/-0.17$)	...
	2958.53	23.77 ($+0.27/-0.22$)	23.07 ($+0.64/-0.40$)
	2962.52	24.14 ($+0.52/-0.35$)	...
	2966.54	24.07 ($+0.57/-0.37$)	23.19 ($+0.21/-0.18$)
	2970.59	24.46 ($+0.41/-0.30$)	23.10 ($+0.22/-0.19$)
	2972.56	24.30 ($+0.28/-0.23$)	23.83 ($+0.88/-0.48$)
	2976.56	24.49 ($+0.90/-0.49$)	23.39 ($+0.26/-0.21$)
	2986.55	>24.81	23.79 ($+0.40/-0.29$)
	2990.54	...	23.80 ($+0.65/-0.40$)
	2994.55	...	24.83 ($+1.68/-0.63$)
2003kp	2936.65	> 24.90	...
	2942.66	24.09 ($+0.26/-0.21$)	24.21 ($+1.18/-0.55$)
	2944.62	23.84 ($+0.35/-0.26$)	23.43 ($+0.20/-0.17$)
	2960.68	...	22.41 ($+0.31/-0.24$)
	2964.68	22.67 (± 0.06)	22.42 (± 0.12)
	2968.67	23.20 ($+0.56/-0.37$)	22.60 (± 0.11)
	2970.66	22.96 (± 0.08)	...
	2970.68	22.92 (± 0.05)	22.53 (± 0.11)
	2972.67	23.07 (± 0.06)	22.47 (± 0.10)
	2974.62	23.14 (± 0.10)	22.58 (± 0.12)
	2976.63	23.21 (± 0.09)	22.55 (± 0.15)
	2988.68	24.24 ($+0.24/-0.20$)	23.51 ($+0.55/-0.36$)
	2992.68	23.80 ($+0.24/-0.19$)	23.84 ($+0.58/-0.38$)
	2996.67	24.30 ($+0.47/-0.33$)	23.83 ($+0.75/-0.44$)
	2998.65	24.41 ($+0.73/-0.43$)	...
2003ku	2934.57	> 24.87	> 25.29
	2940.58	> 24.93	> 24.94
	2958.58	23.06 (± 0.14)	23.03 (± 0.18)

Table 3—Continued

SN name	JD ^b	R_{nat}	I_{nat}
...	2962.59	22.94 (± 0.09)	22.70 (± 0.15)
	2972.59	22.57 (± 0.07)	22.20 (± 0.08)
	2974.55	22.58 (± 0.04)	22.22 (± 0.09)
	2986.59	23.26 (± 0.07)	22.52 (± 0.11)
	2990.57	23.55 (± 0.16)	23.03 (± 0.19)
	2994.58	23.94 (± 0.14)	22.82 (± 0.13)
	3000.59	24.15 (± 0.17)	23.16 (± 0.14)
e510	2942.53	...	24.05 ($+0.32/-0.25$)
	2960.54	...	23.14 ($+0.38/-0.28$)
	2964.54	24.14 ($+0.42/-0.30$)	23.06 ($+0.28/-0.22$)
	2966.62	23.81 ($+0.24/-0.20$)	22.95 ($+0.25/-0.20$)
	2970.57	24.50 ($+1.07/-0.53$)	23.12 ($+0.21/-0.18$)
	2972.53	24.27 ($+0.44/-0.31$)	23.04 ($+0.17/-0.15$)
	2976.54	24.43 ($+0.98/-0.51$)	23.39 ($+0.50/-0.34$)
	2988.56	...	23.70 ($+0.52/-0.35$)
	2994.54	24.64 ($+1.40/-0.59$)	23.89 ($+0.48/-0.33$)
	2996.55	> 25.36	23.90 ($+0.38/-0.28$)
	2998.55	...	24.02 ($+1.24/-0.56$)
	3000.55	...	24.34 ($+0.85/-0.47$)
	2003kv	2936.59	> 23.80
2003kv	2942.60	24.60 ($+0.46/-0.32$)	> 25.00
	2944.57	24.41 ($+1.48/-0.60$)	23.37 ($+0.44/-0.31$)
	2960.61	23.36 ($+0.14/-0.12$)	...
	2964.62	23.47 ($+0.17/-0.15$)	22.86 ($+0.19/-0.16$)
	2968.57	23.63 ($+0.35/-0.27$)	23.35 ($+0.24/-0.20$)
	2970.62	23.48 ($+0.16/-0.14$)	23.08 ($+0.29/-0.23$)
	2970.68	23.73 ($+0.16/-0.14$)	...
	2972.62	23.89 ($+0.30/-0.23$)	22.83 ($+0.17/-0.14$)
	2974.57	23.82 ($+0.15/-0.13$)	23.11 ($+0.20/-0.17$)
	2976.58	23.75 ($+0.23/-0.19$)	23.26 ($+0.27/-0.22$)
	2988.62	24.03 ($+0.25/-0.20$)	23.94 ($+1.03/-0.52$)

Table 3—Continued

SN name	JD ^b	R_{nat}	I_{nat}
...	2992.64	25.11 (+1.41/−0.59)	...
...	2996.62	24.35 (+0.48/−0.33)	24.36 (+0.90/−0.49)
...	2998.61	24.53 (+0.74/−0.43)	23.88 (+0.41/−0.30)
...	3000.61	24.61 (+0.41/−0.30)	24.51 (+0.87/−0.48)
2003lh	2962.62	> 25.54	> 24.57
...	2968.62	23.82 (+0.26/−0.21)	24.63 (+2.58/−0.70)
...	2986.62	22.53 (±0.07)	22.92 (+0.29/−0.23)
...	2990.60	22.71 (±0.06)	22.60 (±0.15)
...	2994.61	22.75 (±0.07)	22.97 (+0.22/−0.18)
...	2996.64	23.00 (±0.10)	22.73 (±0.12)
...	2998.63	22.90 (±0.08)	23.14 (+0.19/−0.16)
...	3000.62	23.17 (±0.08)	23.14 (+0.23/−0.19)
2003le	2962.58	23.81 (+0.28/−0.23)	...
...	2966.58	23.79 (+0.31/−0.24)	> 24.63
...	2972.59	22.98 (±0.08)	23.05 (±0.12)
...	2974.54	22.80 (±0.07)	22.72 (±0.10)
...	2986.58	22.46 (±0.06)	22.29 (±0.08)
...	2990.56	22.55 (±0.07)	22.17 (±0.07)
...	2994.56	22.49 (±0.05)	22.20 (±0.11)
...	2996.58	22.58 (±0.09)	22.44 (±0.11)
...	2998.58	22.59 (±0.05)	22.39 (±0.13)
...	3000.58	22.72 (±0.06)	22.42 (±0.08)
2003ll	2964.66	25.09 (+1.70/−0.63)	> 24.90
...	2968.65	24.42 (+1.23/−0.56)	...
...	2972.67	23.64 (±0.19)	24.09 (+0.59/−0.38)
...	2974.61	23.53 (±0.19)	22.47 (±0.08)
...	2976.62	23.10 (±0.13)	22.58 (±0.13)
...	2988.66	23.14 (±0.13)	23.60 (+0.81/−0.46)
...	2992.67	23.24 (±0.08)	24.05 (+0.75/−0.44)
...	2996.66	24.88 (+0.27/−0.21)	...
2003li	2958.65	> 25.64	> 25.38

Table 3—Continued

SN name	JD ^b	R_{nat}	I_{nat}
...	2962.66	> 25.57	> 24.87
...	2966.66	24.76 (+0.34/−0.26)	23.79 (+0.62/−0.39)
...	2986.66	22.75 (±0.06)	22.49 (±0.08)
...	2990.64	22.69 (±0.07)	22.76 (±0.11)
...	2994.64	...	22.67 (±0.09)
...	2996.66	23.08 (±0.11)	22.88 (±0.09)
...	2998.65	23.15 (±0.09)	22.78 (±0.11)

^aThe values given are natural system “Vega” magnitudes. See text for more details. If a value is given as “> some number,” it is a 1σ upper limit. The values in parentheses are 1σ lower and upper error bars.

^bJulian Date minus 2,450,000.

Table 4. MLCS Fits to Seven ESSENCE SNe Ia^a

SN	z	$t(B_{max})$ ^b	MLCS Δ	A_V^c (mag)	$m - M$ (mag)	$\Delta(m - M)$
2003jo	0.53	2934.93	−0.40(0.09)	0.44(0.10)	42.67(0.20)	+0.20(0.20)
2003kp	0.64	2962.80	−0.40(0.11)	0.16(0.13)	43.11(0.20)	+0.14(0.20)
2003kv	0.78	2963.70	−0.40(0.15)	0.30(0.21)	43.46(0.27)	−0.05(0.27)
2003lh	0.54	2984.76	−0.40(0.14)	0.06(0.14)	43.19(0.20)	+0.67(0.20)
2003le	0.56	2985.90	−0.40(0.13)	0.13(0.12)	42.76(0.21)	+0.15(0.21)
2003ll	0.60	2979.79	+0.52(0.28)	0.28(0.26)	42.08(0.52)	−0.72(0.52)
2003li	0.54	2986.11	−0.40(0.14)	0.24(0.14)	43.03(0.24)	+0.51(0.24)

^aUsing Version 3 of the Multi-color Light Curve Shape method, MLCS2k2 (Jha, Riess, & Kirshner 2005). These results allow no extrapolation beyond the training set. The last column contains the arithmetic differences of the derived distance moduli and those expected in an empty universe ($\Omega_M = 0.0, \Omega_\Lambda = 0.0$).

^bTime of B -band maximum. Julian Date minus 2,450,000.

^cHost-galaxy extinction. The Galactic extinction has been subtracted, using the color excesses given in Table 1.

Table 5. BATM Fits to Seven ESSENCE SNe Ia^a

SN	z	A_V^b (mag)	$m - M$ (mag)	$\Delta(m - M)$
2003jo	0.53	0.21(0.25)	42.89(0.31)	+0.44(0.31)
2003kp	0.64	0.01(0.11)	43.11(0.19)	+0.16(0.19)
2003kv	0.78	0.02(0.12)	44.12(0.21)	+0.62(0.21)
2003lh	0.54	0.03(0.13)	43.20(0.22)	+0.70(0.22)
2003le	0.56	0.03(0.13)	42.86(0.15)	+0.26(0.15)
2003ll	0.60	0.60(0.32)	43.03(0.53)	+0.25(0.53)
2003li	0.54	0.17(0.26)	43.03(0.35)	+0.53(0.35)

^aUsing the Bayesian Adapted Template Method (Tonry et al. 2003, and Tonry et al. 2005, in preparation). The last column contains the arithmetic differences of the derived distance moduli and those expected in an empty universe ($\Omega_M = 0.0, \Omega_\Lambda = 0.0$).

^bHost-galaxy extinction. The Galactic extinction has been subtracted, using the color excesses given in Table 1.

Table 6. $\Delta m_{15}(B)$ Fits to Seven ESSENCE SNe Ia^a

SN	z	χ_ν^2	$\Delta m_{15}(B)$	$E(B - V)_{host}$	$m - M$ (mag)	$\Delta(m - M)$
2003jo	0.53	0.57	0.83(0.04)	0.09(0.02)	42.62(0.19)	+0.17(0.19)
2003kp	0.64	1.80	0.88(0.02)	0.01(0.02)	43.09(0.18)	+0.14(0.18)
2003kv	0.78	0.56	0.92(0.02)	0.06(0.03)	43.54(0.20)	+0.04(0.20)
2003lh	0.54	2.49	0.87(0.03)	0.07(0.06)	42.63(0.21)	+0.13(0.21)
2003le	0.56	1.56	0.84(0.04)	0.06(0.02)	42.51(0.19)	-0.09(0.19)
2003ll	0.60	3.35	1.30(0.02)	0.08(0.07)	42.62(0.29)	-0.16(0.29)
2003li	0.54	2.90	0.83(0.04)	0.11(0.04)	42.68(0.21)	+0.18(0.21)

^aThese values are from the light-curve fits using the $\Delta m_{15}(B)$ method of Prieto, Rest, & Suntzeff (2005). The last column contains the arithmetic differences of the derived distance moduli and those expected in an empty universe ($\Omega_M = 0.0, \Omega_\Lambda = 0.0$).

Fig. 1.— $10'' \times 10''$ images of the nine ESSENCE SNe discussed in this paper. Each was taken through the F775W filter of *HST*/ACS. The total integration times were 3700 s (e510), 5100 s (SN 2003jo), and 4400 s (the other seven SNe).

Fig. 2.— Difference of observed R and I photometry of stars near six of our SNe compared to R and I magnitudes derived from SDSS data using the transformations of Smith et al. (2002). Δ is in the sense “observed values from CTIO 0.9-m photometry directly tied to Landolt (1992) standards” *minus* “values derived from SDSS photometry.” (In the color version of this plot different colors correspond to different ESSENCE fields.)

Fig. 3.— Observed-frame natural system photometry of the nine SNe discussed in this paper. Symbols: R band = (blue) squares; *HST*/ACS F625W filter = (yellow) dots. Downward pointing triangles are 1σ upper limits in R .

Fig. 4.— Observed-frame natural system photometry of the nine SNe discussed in this paper. Symbols: I band = (green) squares; *HST*/ACS F775W filter = (orange) dots; *HST*/ACS F850LP filter = (red) upward-pointing triangles. Downward-pointing triangles are 1σ upper limits in I .

Fig. 5.— Light-curve fits in the rest-frame bands. All these fits used the $\Delta m_{15}(B)$ method of Prieto, Rest, & Suntzeff (2005). The symbols are the same as in Figures 3 and 4, with one exception. In Fig. 5f the diamond-shaped symbols correspond to photometry of SN 2003 ll originally obtained in the I -band.

Fig. 6.— K-corrected, extinction-corrected I -band data of SN 2003lh along with I -band data of the slow decliners SN 1999aa (Krisciunas et al. 2000; Jha 2002) and SN 1991T (Lira et al. 1998), adjusted in magnitude space to the brightness of SN 2003lh using the $\Delta m_{15}(B)$ solution in Table 6 and using $M_I(\text{max}) = -19.32$ mag (on an $H_0 = 65 \text{ km s}^{-1} \text{ Mpc}^{-1}$ scale) for the slowest decliners studied by Nobili et al. (2005).

Fig. 7.— Residuals of MLCS2k2 light-curve fits. Here we assume no limits on the possible values of the MLCS parameter Δ . The CTIO 4-m data are represented by squares, while the *HST* data are represented by triangles. Differential magnitude here is in the sense of “rest-frame extinction-corrected, K-corrected data value” *minus* “interpolated light-curve fit value.”

Fig. 8.— Same as for Figure 7, but residuals of data and $\Delta m_{15}(B)$ light-curve fits.

Fig. 9.— Composite spectrum of the six slowly declining ESSENCE objects. SNe 1992A and 1999aa (at ages of 3 and 5 d past maximum brightness for SN 1992A and 3 d past maximum brightness for SN 1999aa) are plotted for comparison. The bars at the bottom of the plot

show the wavelength range of each object’s spectrum used for the composite spectrum.

Fig. 10.— Histogram of decline-rate values [$\Delta m_{15}(B)$] for 107 nearby SNe Ia (Gallagher et al. 2005) and values for seven ESSENCE SNe from Table 6. Clearly, most of the ESSENCE objects have slow decline rates compared to the local sample.

Fig. 11.— Typical pre-maximum detection window for SNe Ia discovered with the CTIO 4-m telescope. We assume $0.9''$ seeing, $S/N \gtrsim 10$, and a detection threshold of $R = 23.0$ mag. ESSENCE fields are typically imaged (or “visited”) every 4 days. Under typical seeing conditions, most of the SNe Ia discovered with $z \gtrsim 0.60$ will have slow decline rates.

Fig. 12.— Spectra of two nearby SNe Ia, redshifted by the indicated amounts, and superimposed on the spectrum of SN 2003ku. “Normalized flux” means that a pseudo-continuum has been subtracted from the spectra.

Fig. 13.— Differential Hubble diagram for SNe Ia. For each object we plot the distance modulus derived from the light curves minus the distance modulus in an empty universe vs. the redshift. See Table 4. The yellow dots are the “gold” data set of 157 SNe Ia from Riess et al. (2004). The ESSENCE SNe Ia are represented by larger squares. SN 2003ll is represented by a red square having a distance modulus 0.72 mag *brighter* than the empty-universe model. The dashed line is the concordance model ($\Omega_M = 0.3, \Omega_\Lambda = 0.7$). The solid horizontal line corresponds to distance moduli in the empty-universe model. The dotted line corresponds to an open universe with $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.0$.

Fig. 14.— Constraints on the dark energy (Ω_Λ) using the nearby SNe Ia of the “gold” set of Riess et al. (2004), plus seven ESSENCE SNe Ia. We show sets of 1σ , 2σ , and 3σ contours. The solid lines use a constraint that MLCS $\Delta \geq -0.40$. The dashed contours assume no constraints on MLCS Δ .

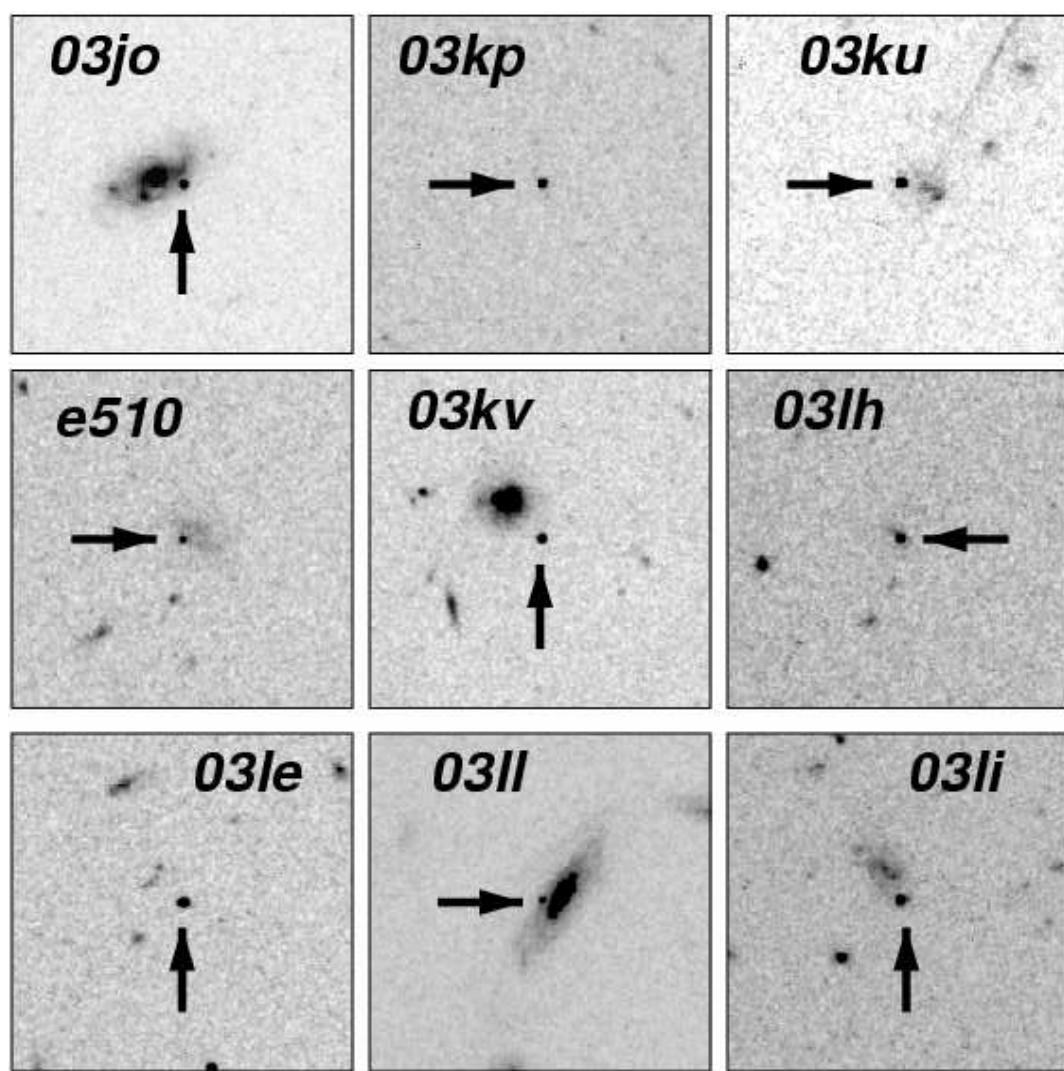
Fig. 15.— Constraints on the dark energy (Ω_Λ) using the entire gold set of 157 SNe Ia from Riess et al. (2004), plus 7 ESSENCE SNe Ia. We show sets of 1σ , 2σ , and 3σ contours. The thin lines come from the SN constraints only and assume the prior constraint that MLCS $\Delta \geq -0.40$. The thick lines include a matter density constraint of $\Omega_M = 0.233 \pm 0.030$, in accord with the results of Cole et al. (2005) and $H_0 = 72 \text{ km s}^{-1}\text{Mpc}^{-1}$ (Freedman et al. 2001). The dashed contours include a different matter constraint, that of Eisenstein et al. (2005), namely $\Omega_M = 0.273 \pm 0.025 + 0.137\Omega_K$. We assume $w = -1$.

Fig. 16.— Constraints on the equation-of-state parameter w using the Riess et al. (2004) gold set of 157 SNe Ia, plus 7 ESSENCE SNe Ia. We show sets of 1σ , 2σ , and 3σ contours. The thin lines come from the SN constraints only, while the thick lines are constrained by

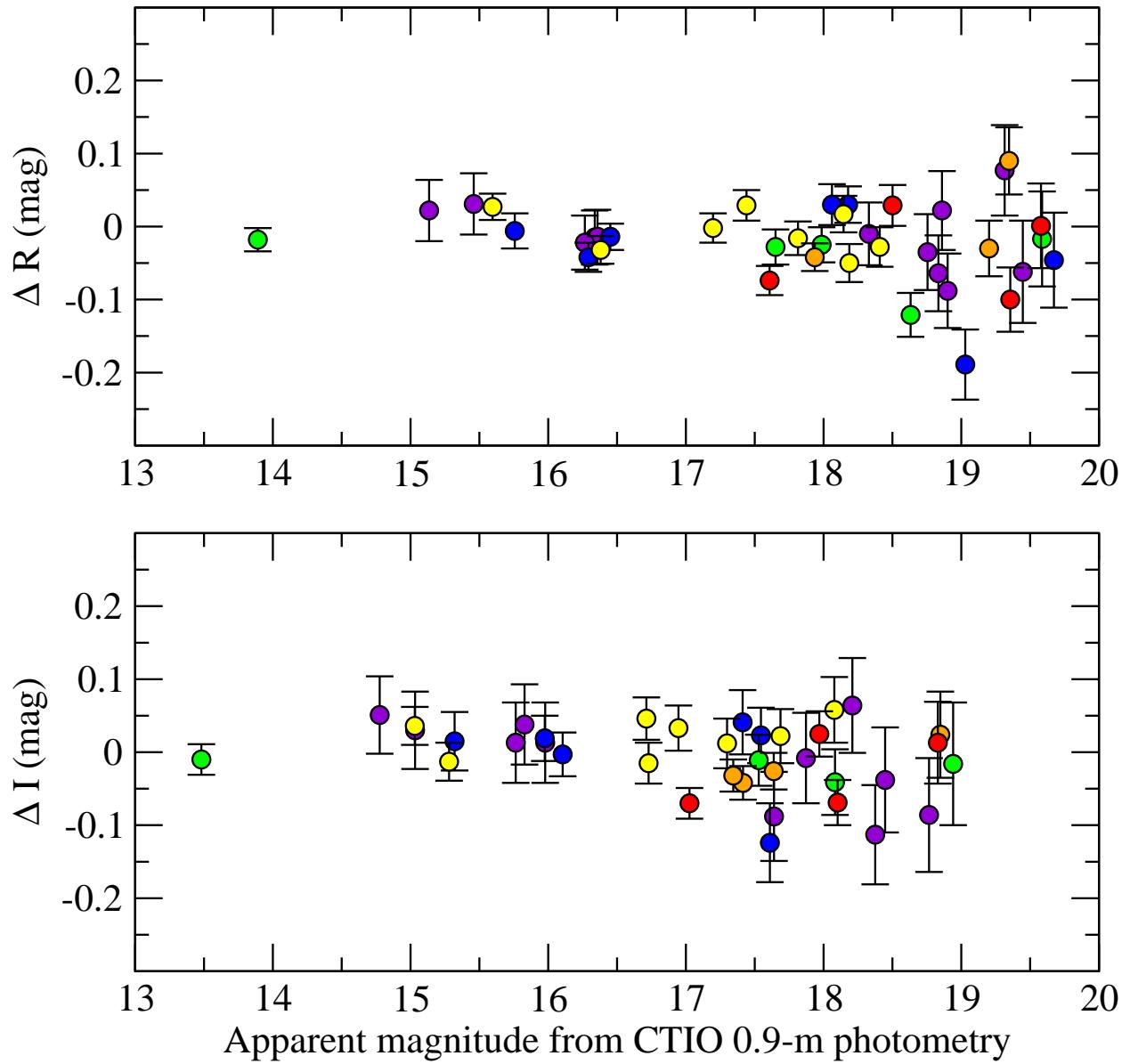
SNe plus a matter density of $\Omega_M = 0.233 \pm 0.030$ (Cole et al. 2005), which implies $H_0 = 72 \text{ km s}^{-1}\text{Mpc}^{-1}$ (Freedman et al. 2001). The dashed contours show the effect of using no prior constraint on the allowable range of MLCS Δ .

Fig. 17.— The throughput of the F625W, F775W, and F850LP filters as used with the *HST/ACS*. The curves include the filter transmission functions multiplied by the quantum efficiency as a function of wavelength.

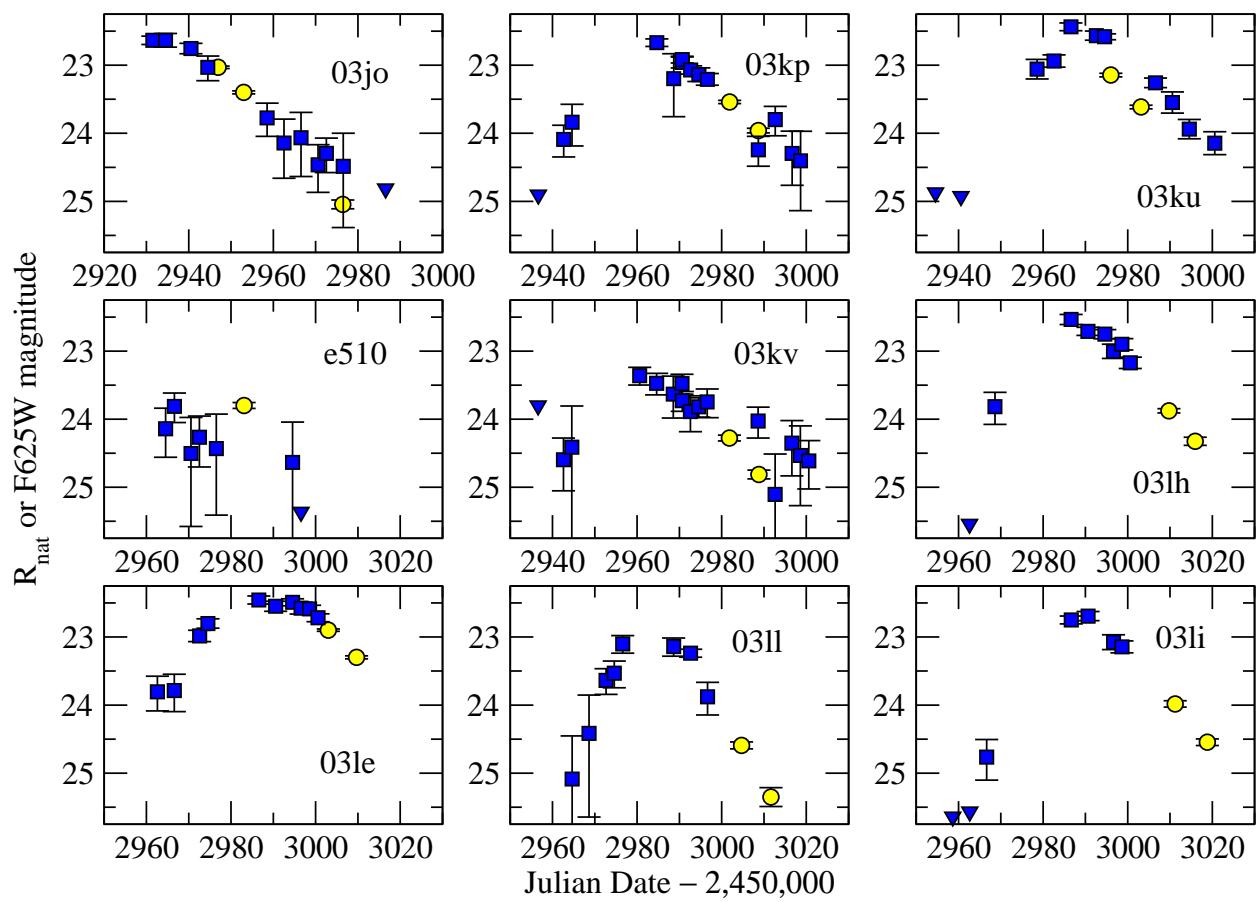
Fig. 18.— Effective filter transmission curves of the *R* and *I* filters used for ground-based photometry with the CTIO 4-m telescope. These curves show the fractional transmission as a function of wavelength. We included the filter transmission functions determined in the laboratory, one aluminum reflection for the effect of the primary mirror, the quantum efficiency of the CCDs, the atmospheric extinction, and the major telluric absorption lines. To obtain the curves used by Bessell (1990) one must multiply these curves by the wavelength.



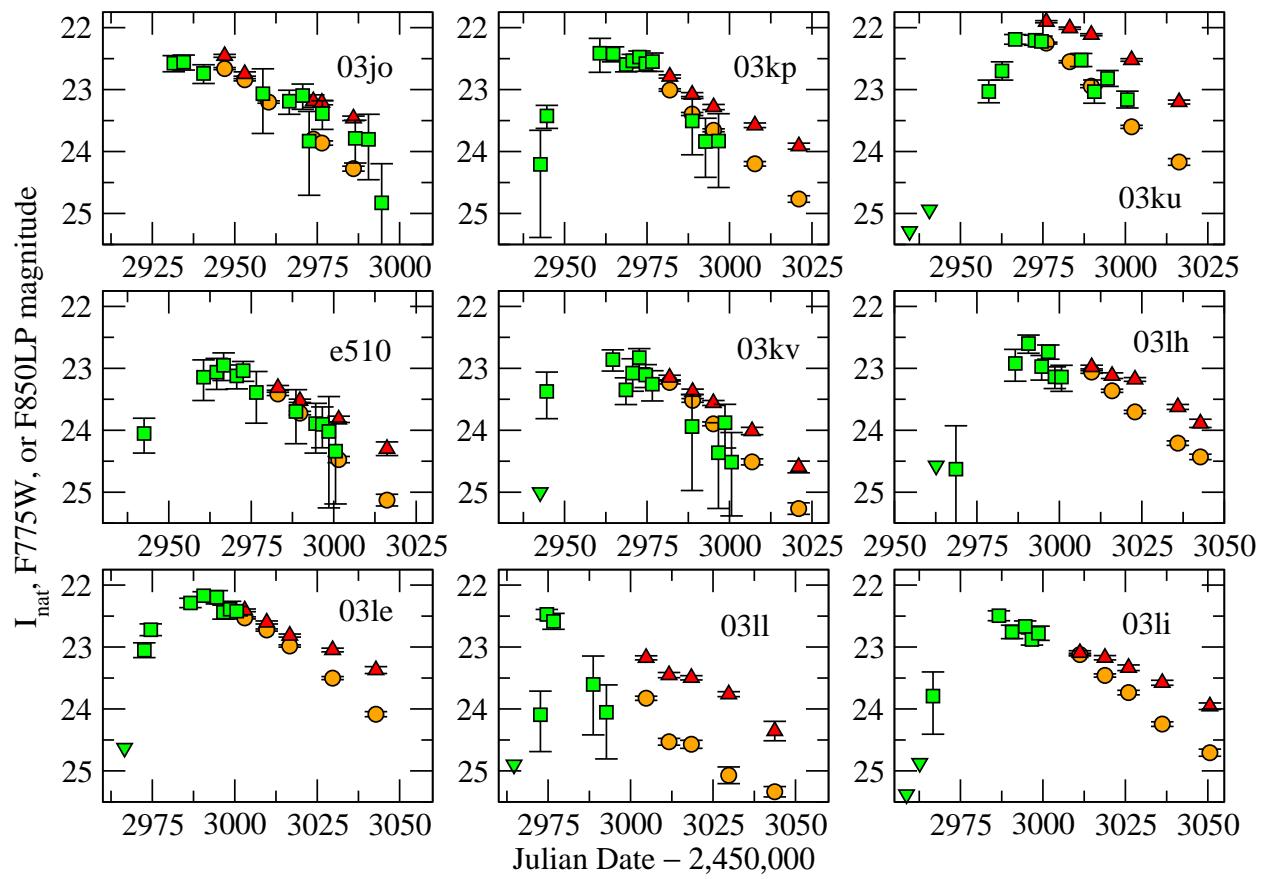
Krisciunas *et al.* Fig. 1



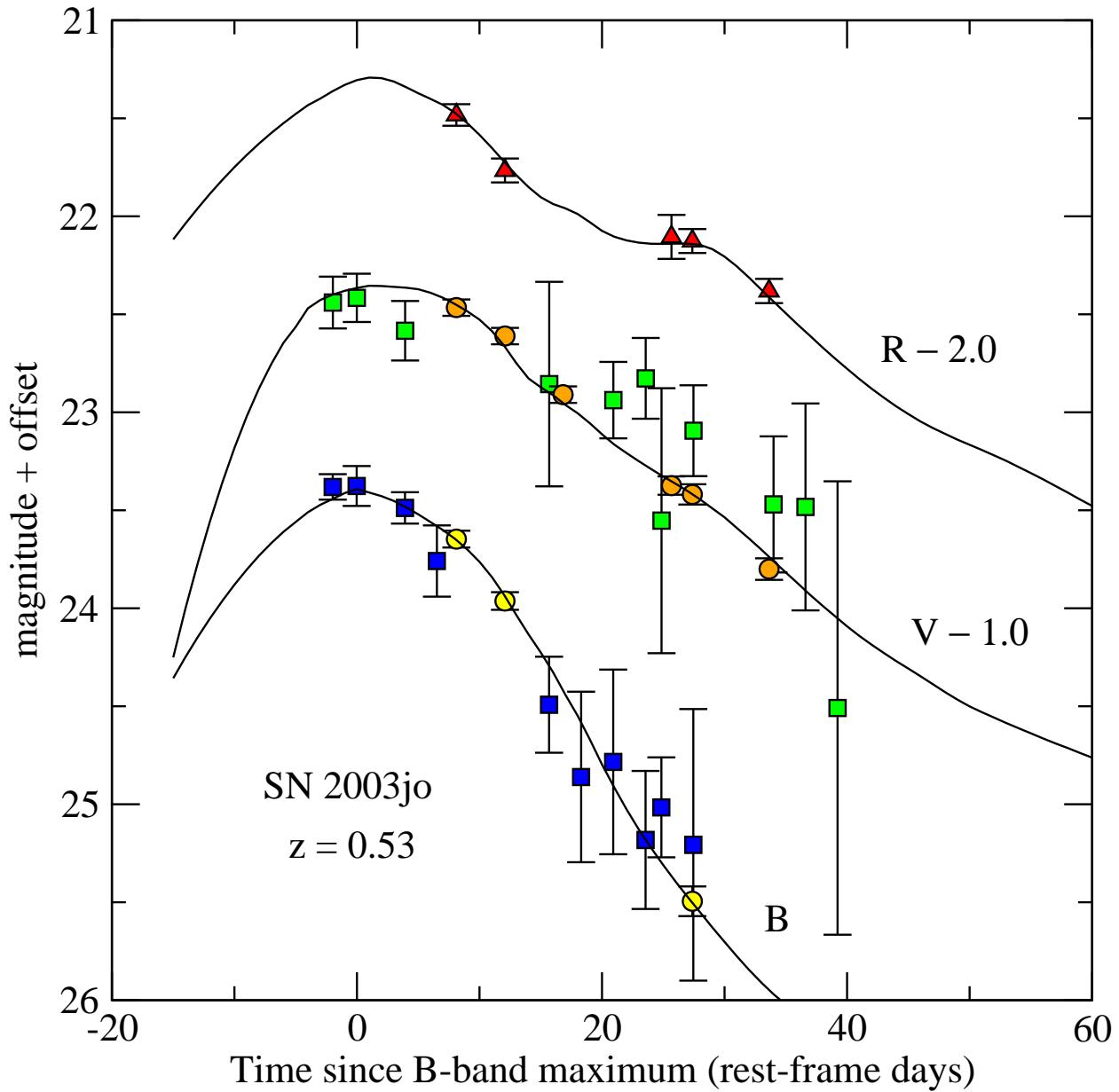
Krisciunas *et al.* Fig. 2



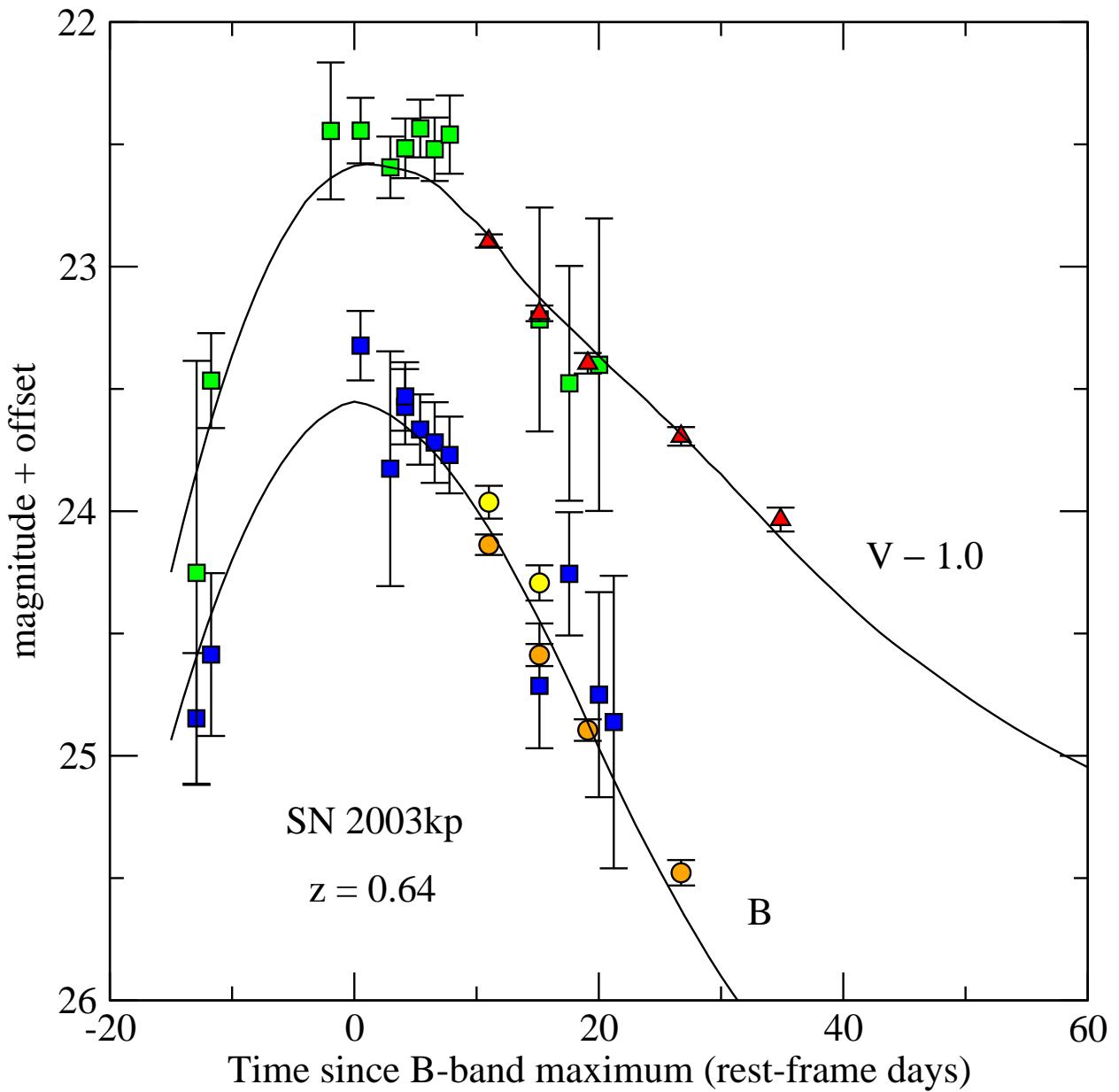
Krisciunas *et al.* Fig. 3



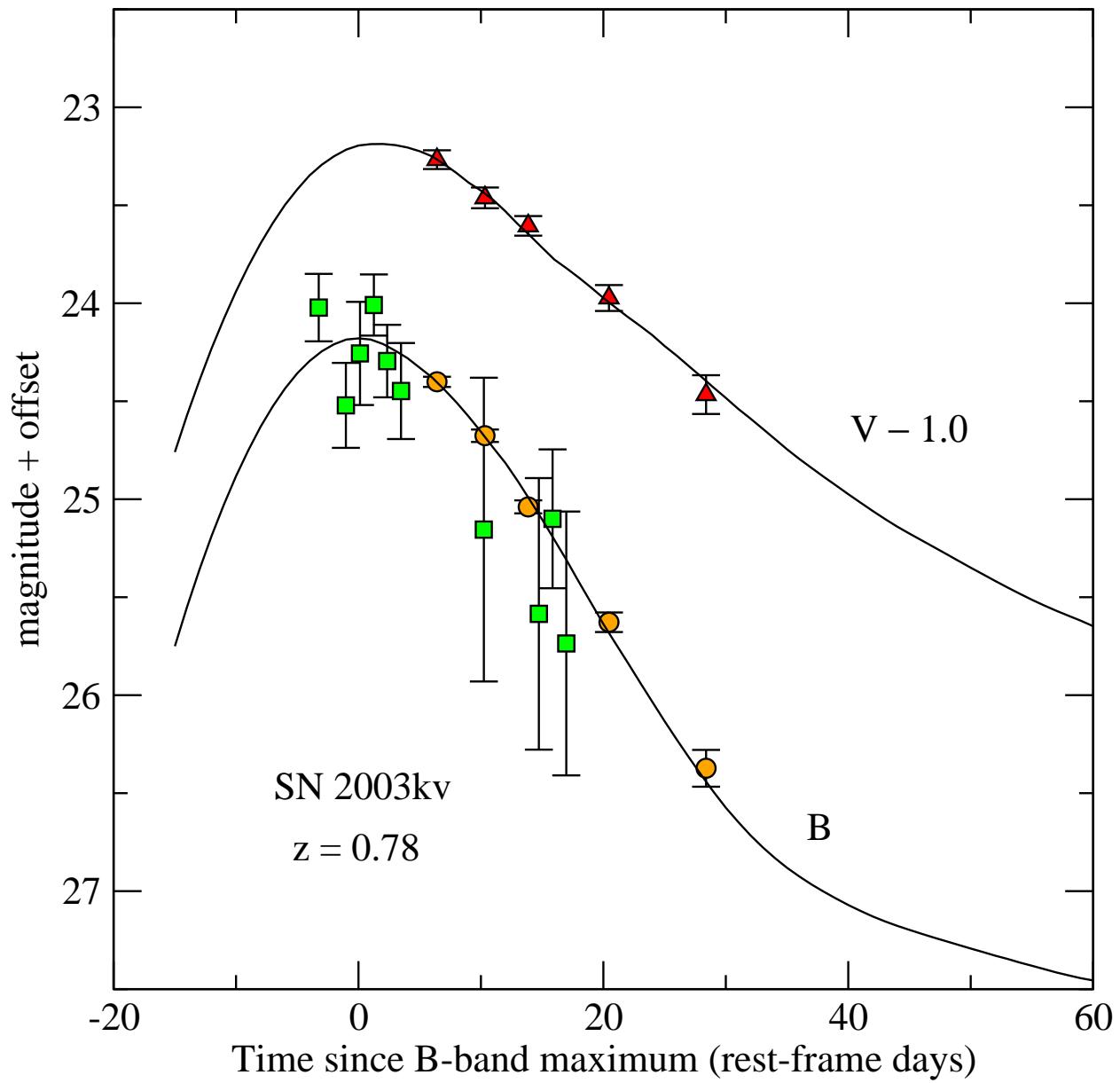
Krisciunas *et al.* Fig. 4



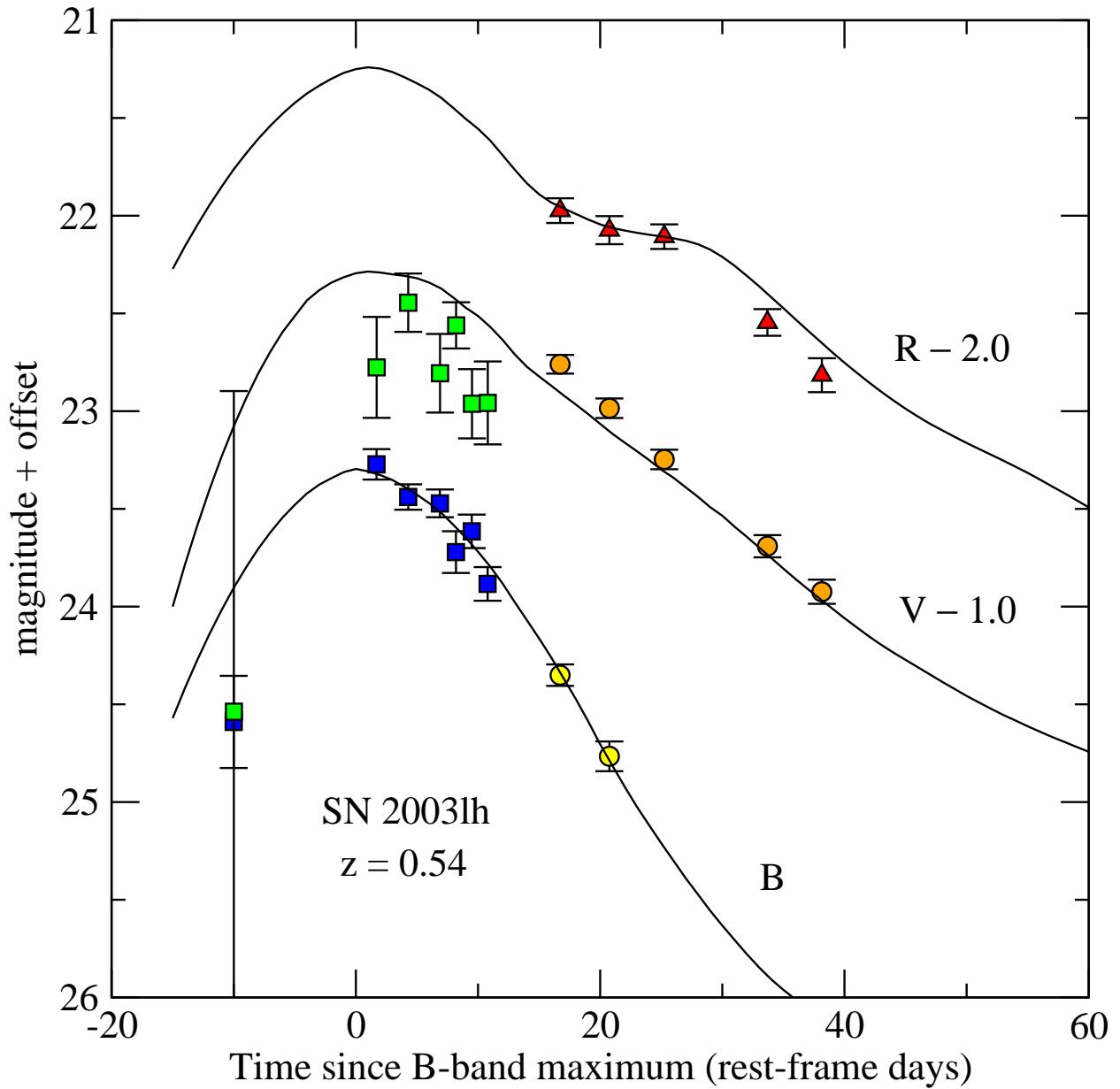
Krisciunas *et al.* Fig. 5a



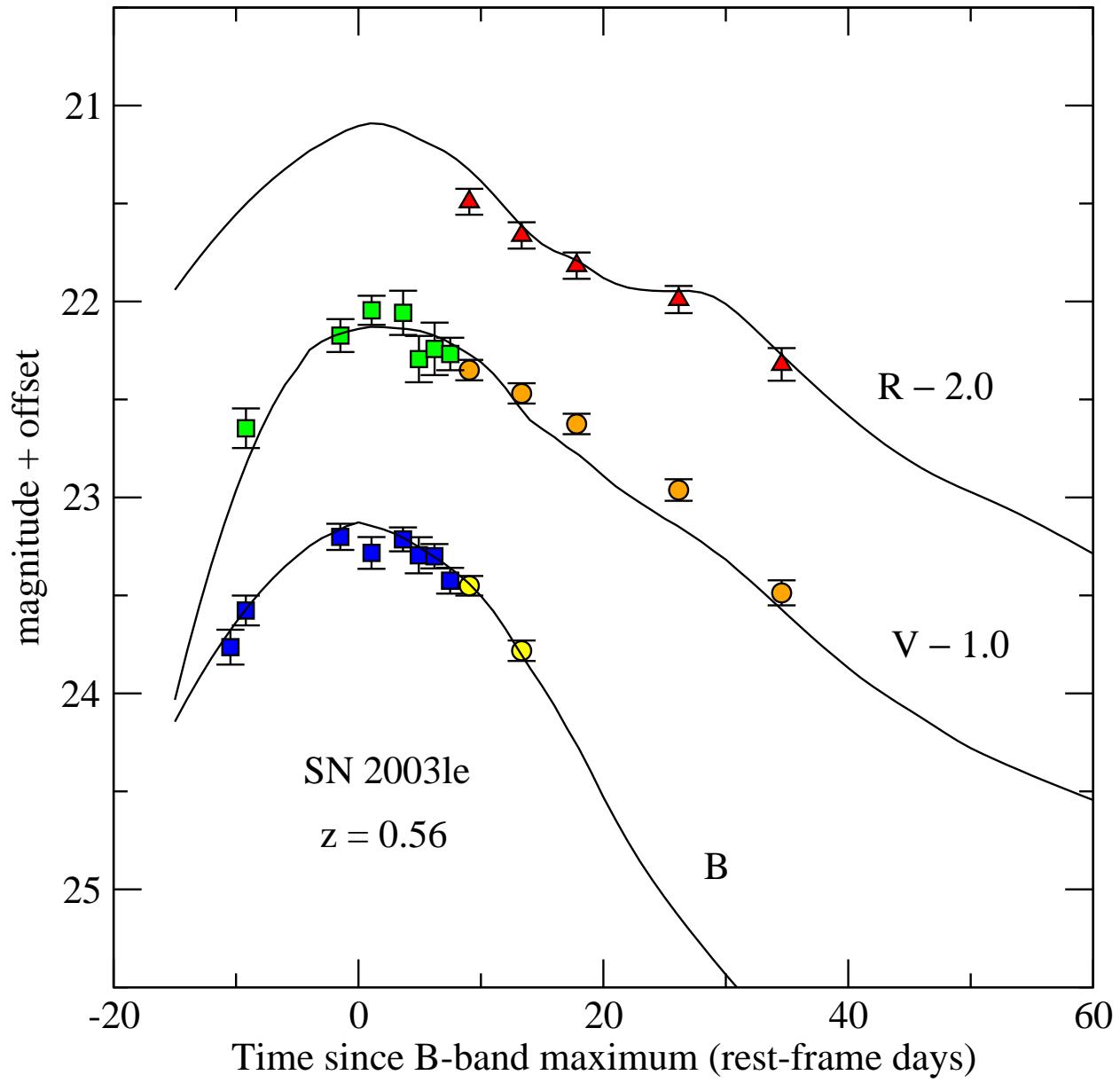
Krisciunas *et al.* Fig. 5b



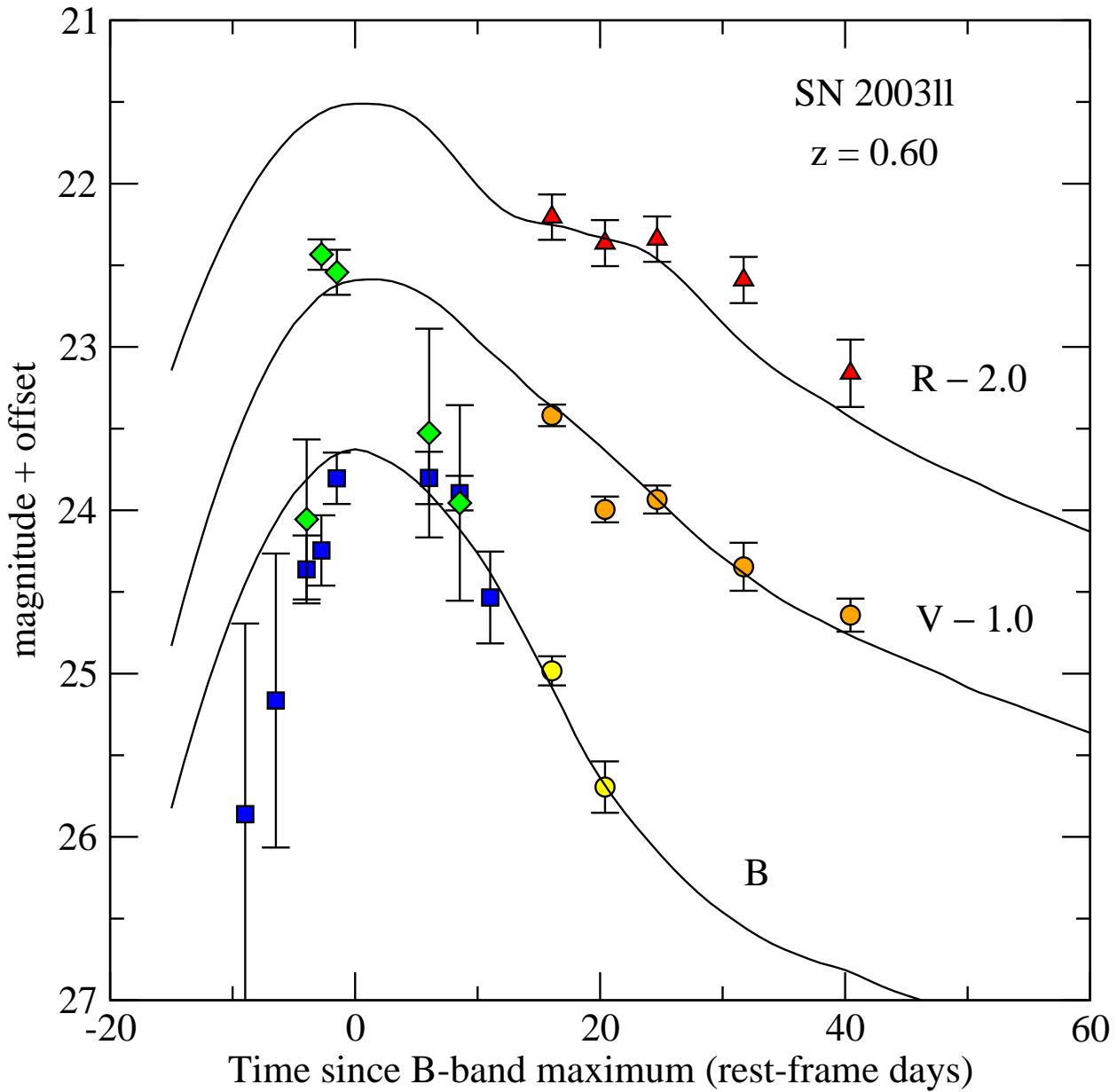
Krisciunas *et al.* Fig. 5c



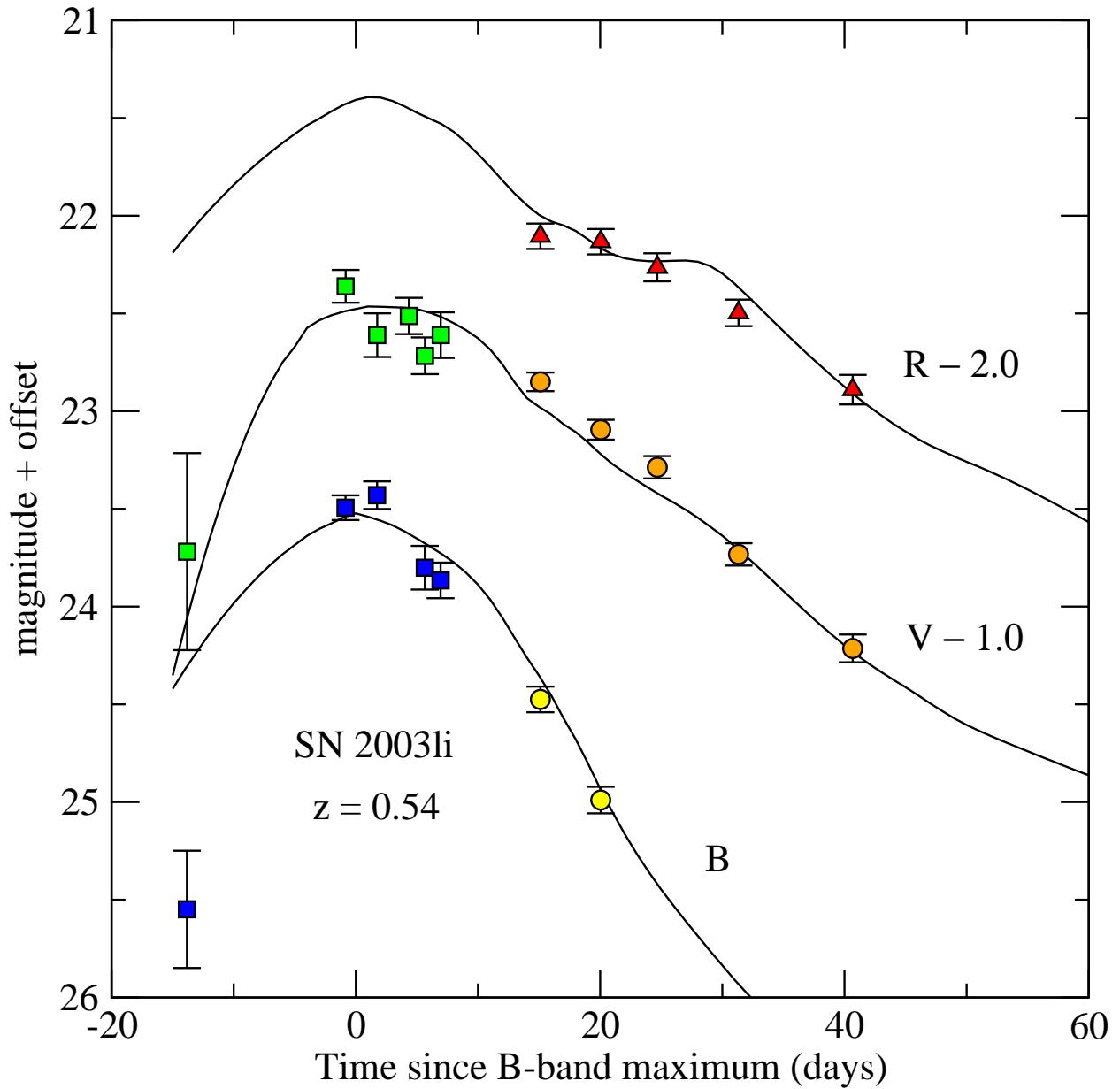
Krisciunas *et al.* Fig. 5d



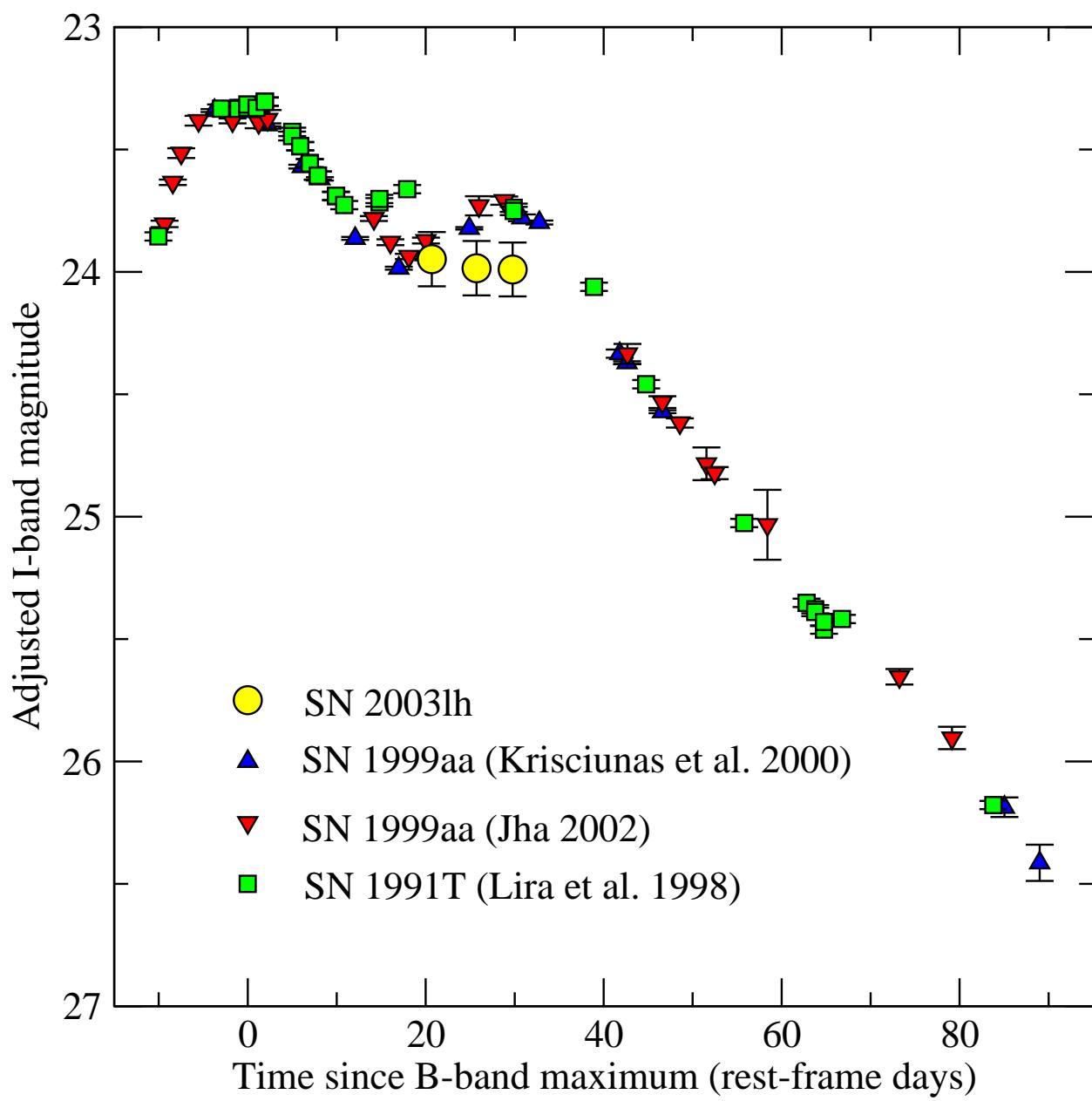
Krisciunas *et al.* Fig. 5e



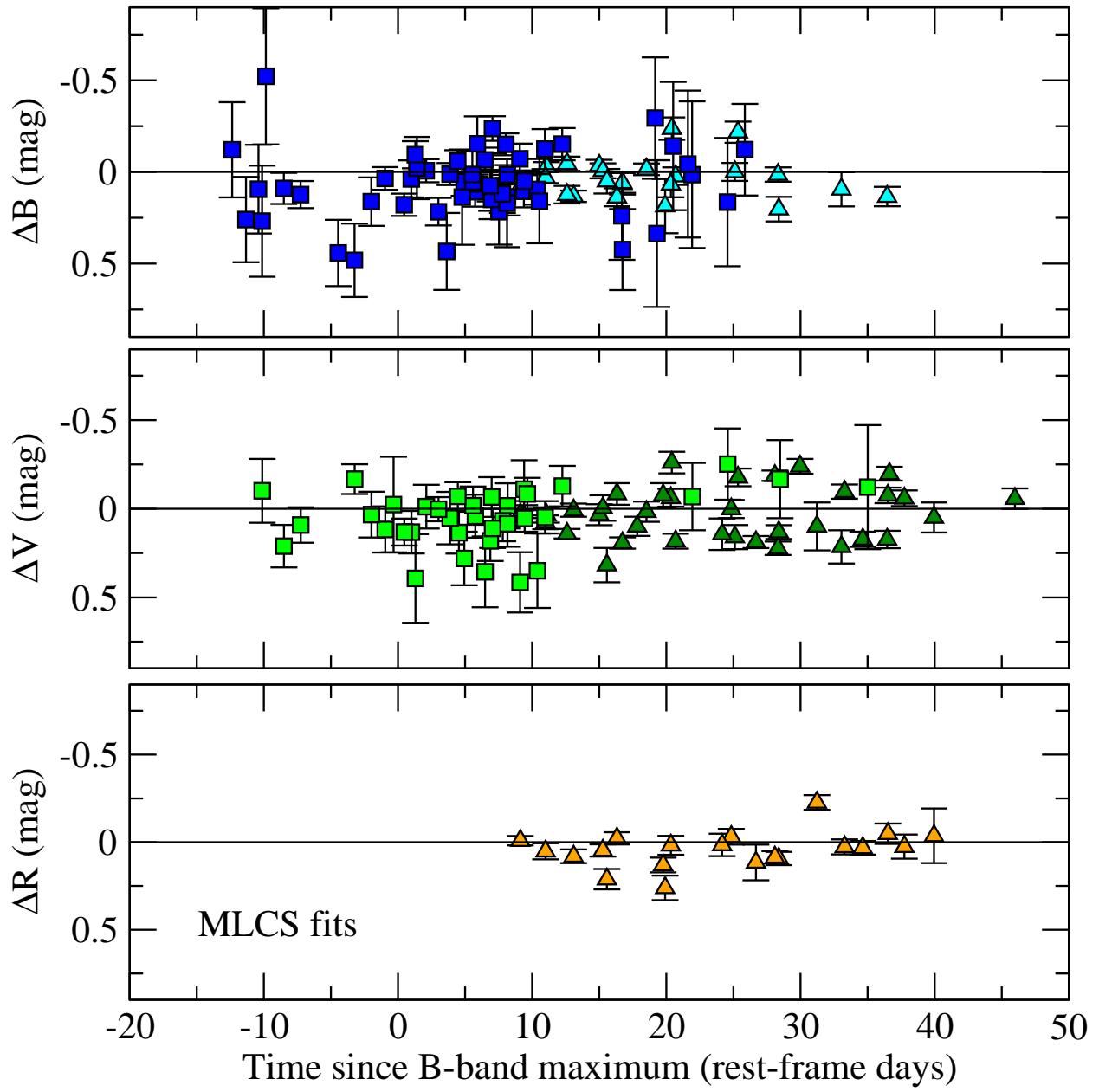
Krisciunas *et al.* Fig. 5f



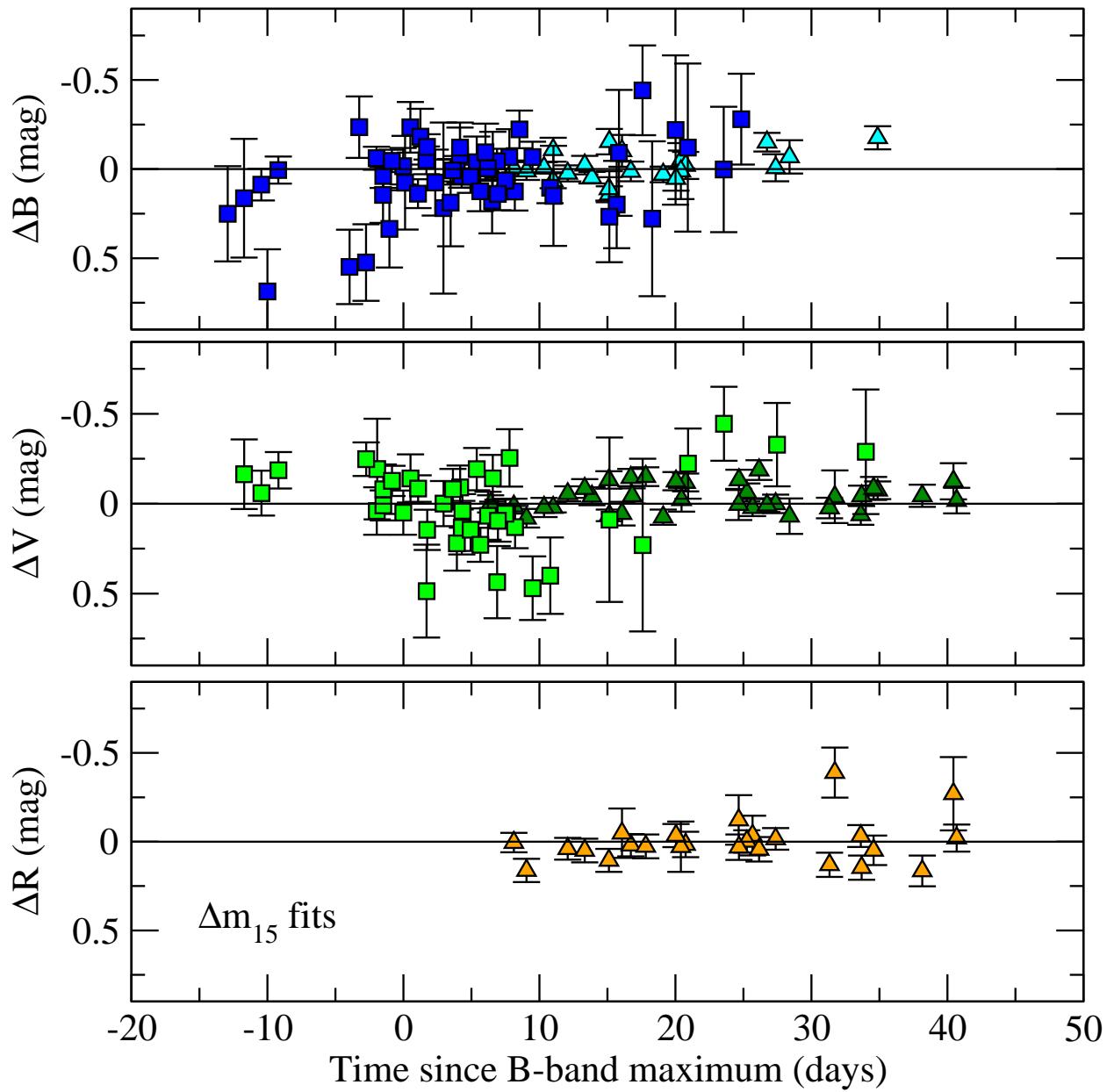
Krisciunas *et al.* Fig. 5g



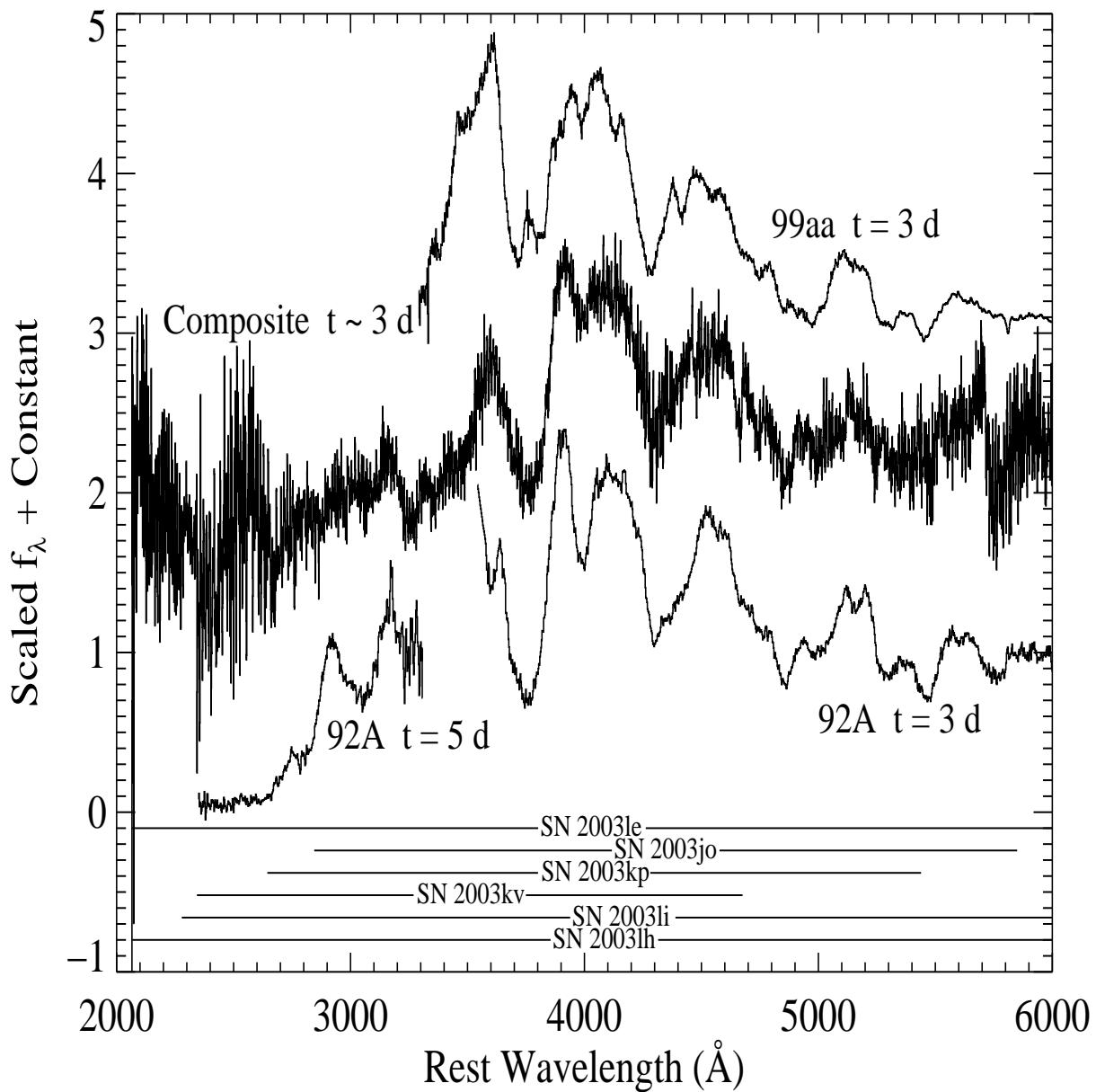
Krisciunas *et al.* Fig. 6



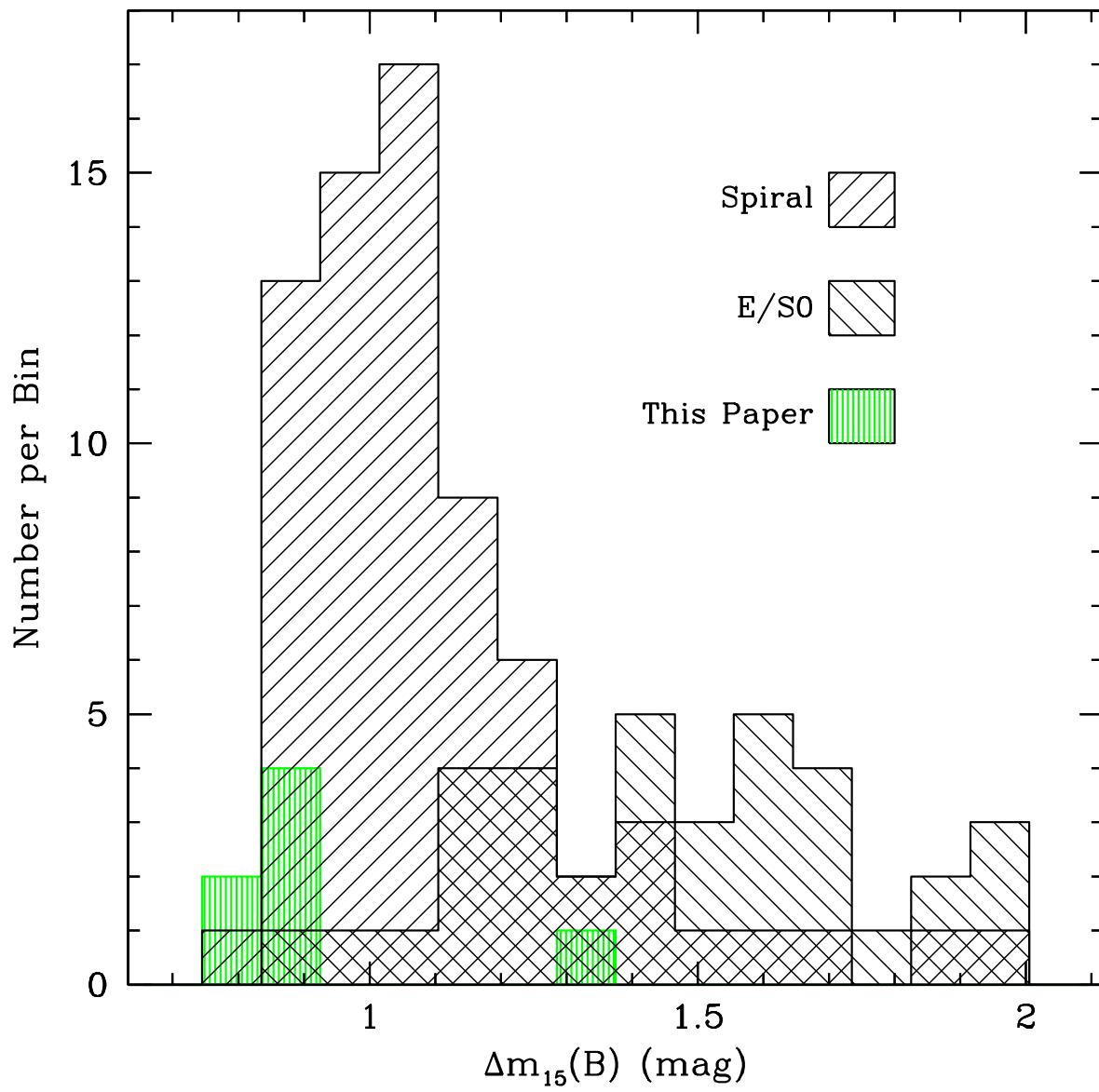
Krisciunas *et al.* Fig. 7



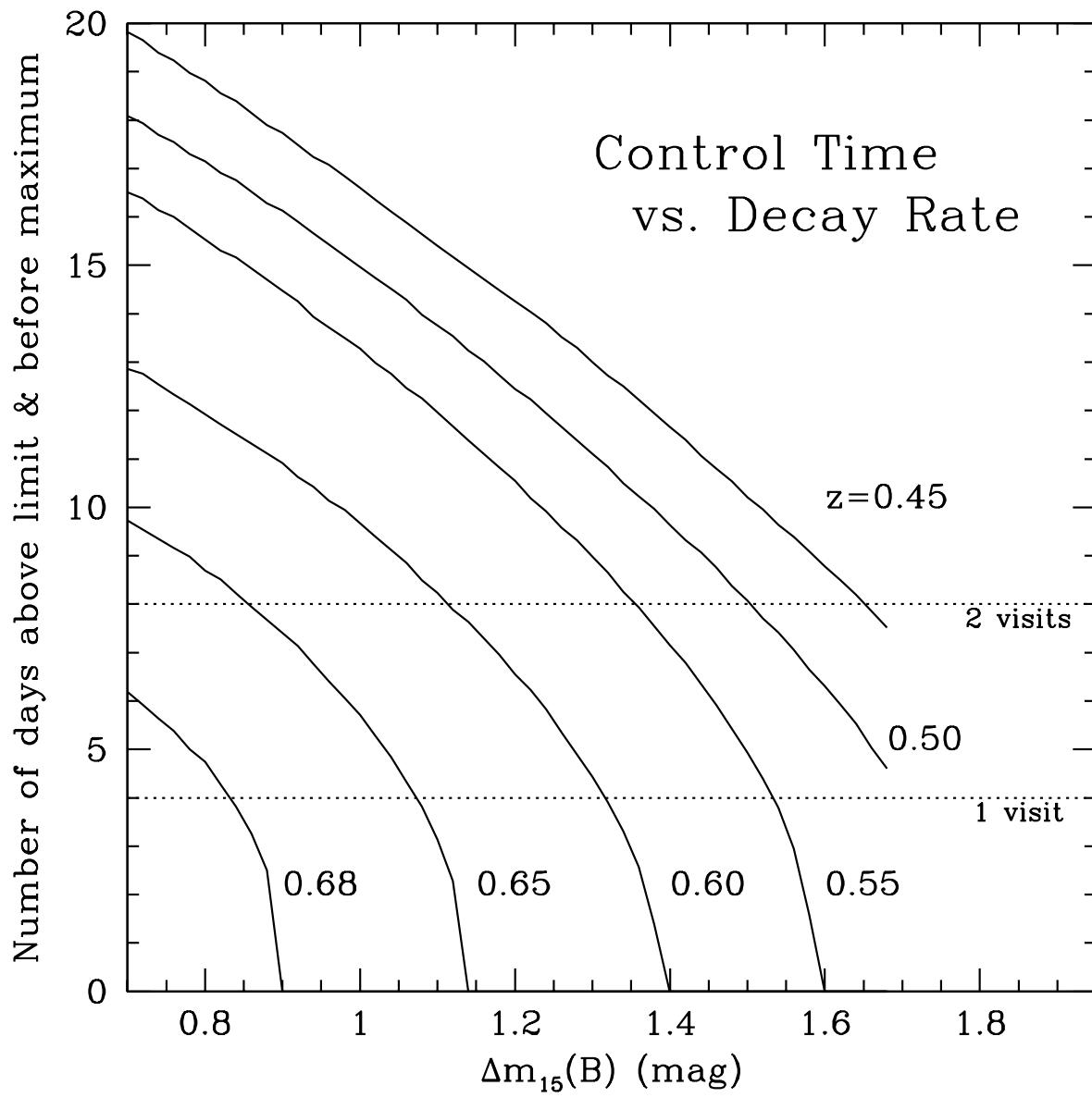
Krisciunas *et al.* Fig. 8



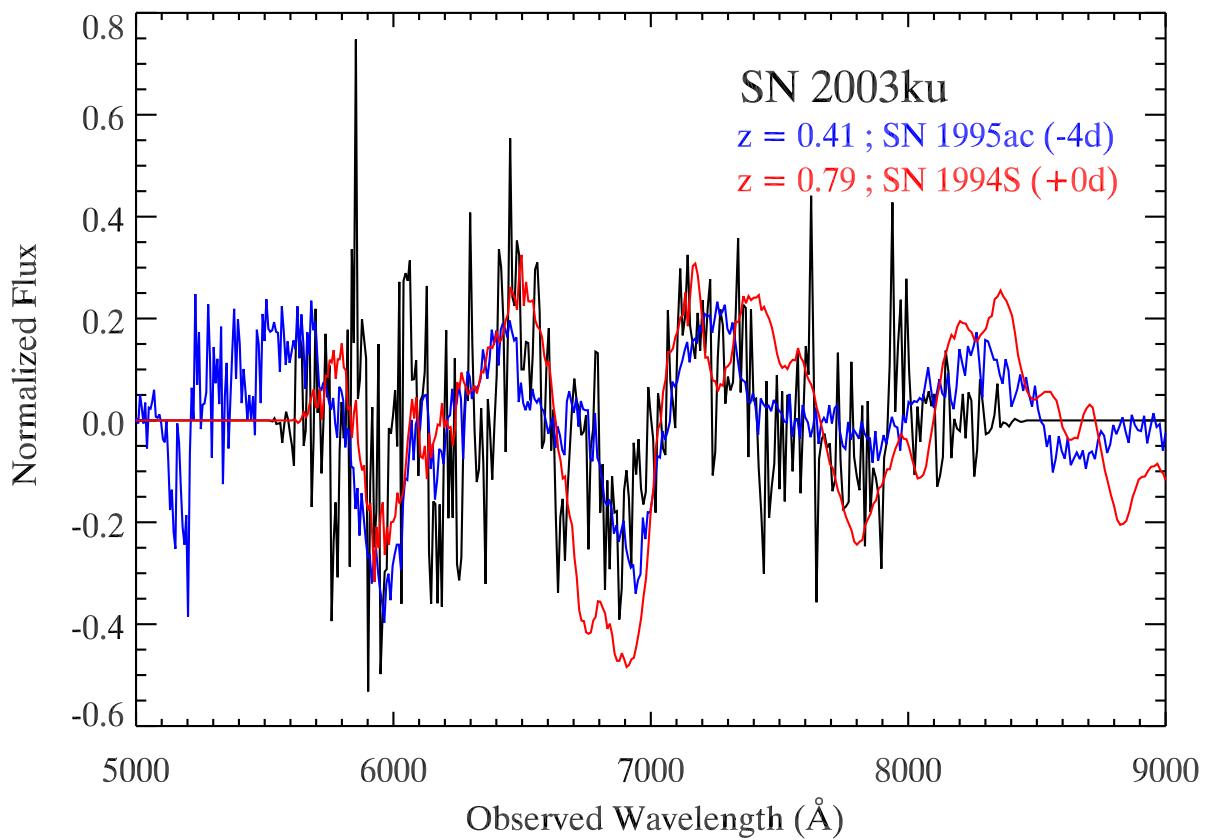
Krisciunas *et al.* Fig. 9



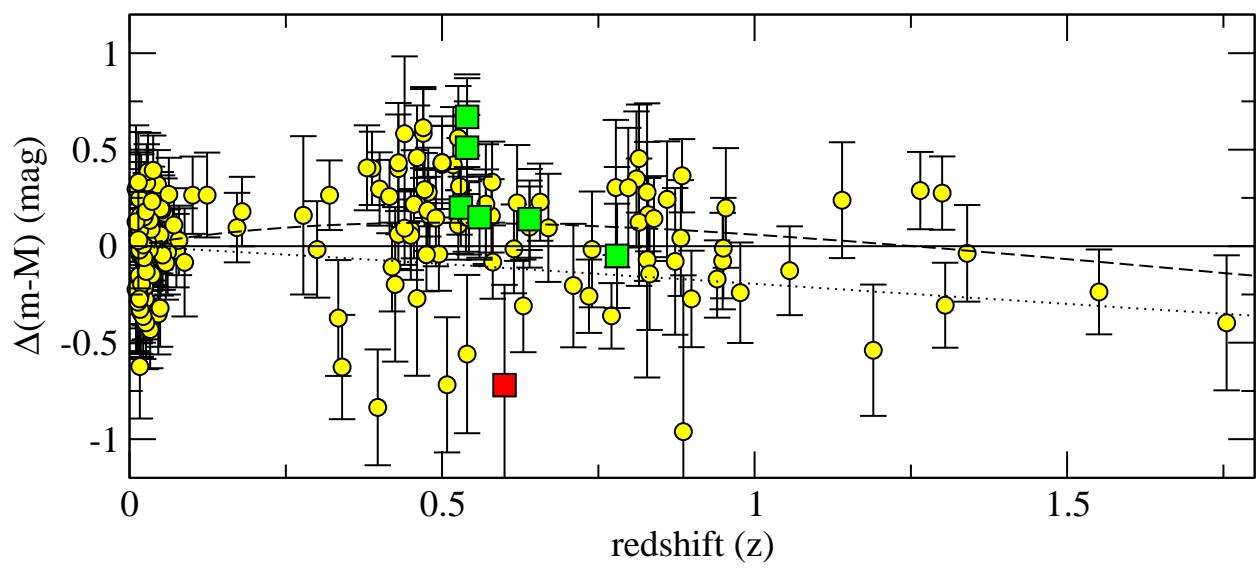
Krisciunas *et al.* Fig. 10



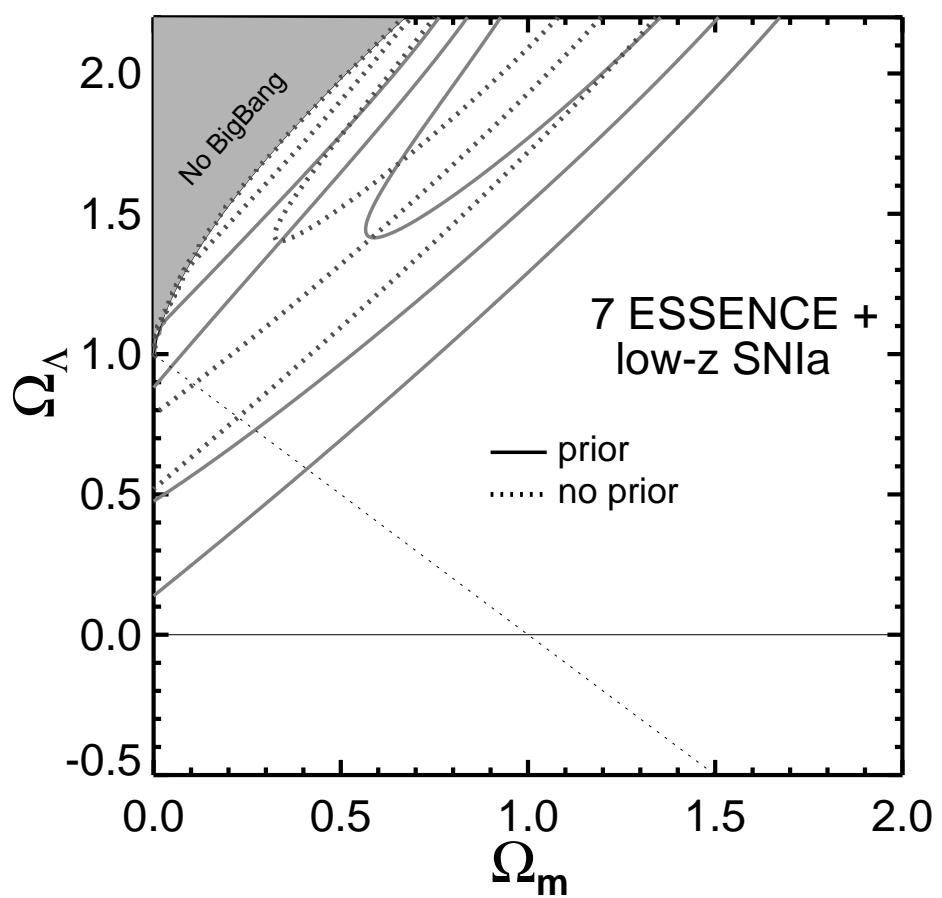
Krisciunas *et al.* Fig. 11



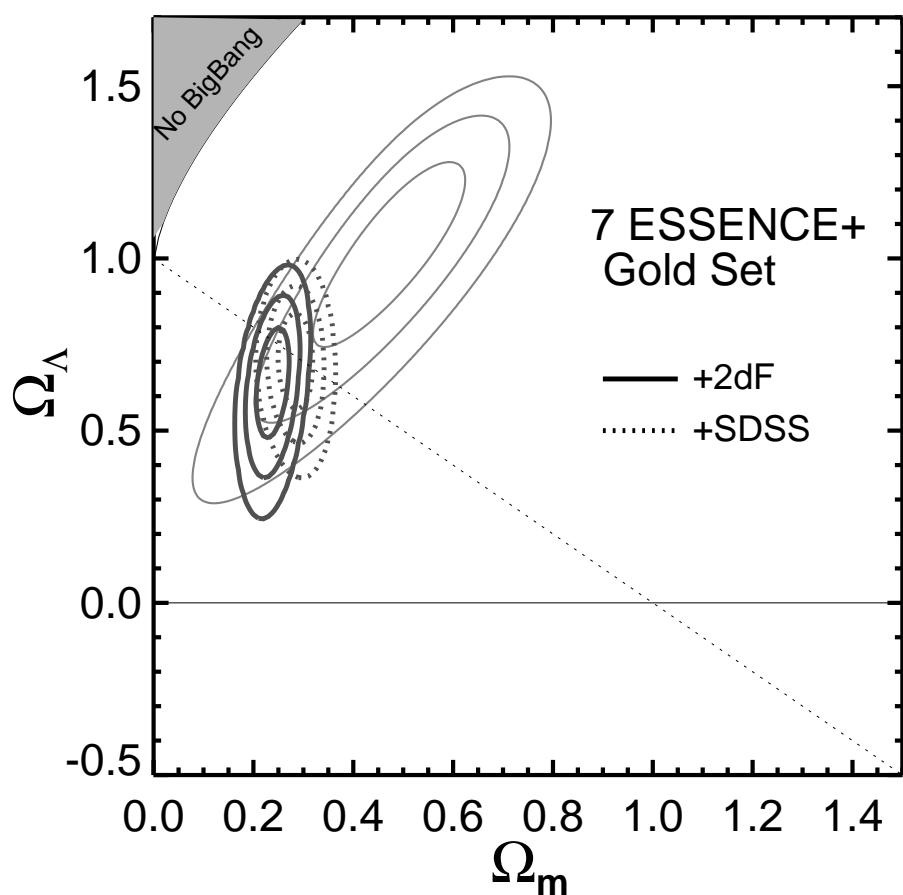
Krisciunas *et al.* Fig. 12



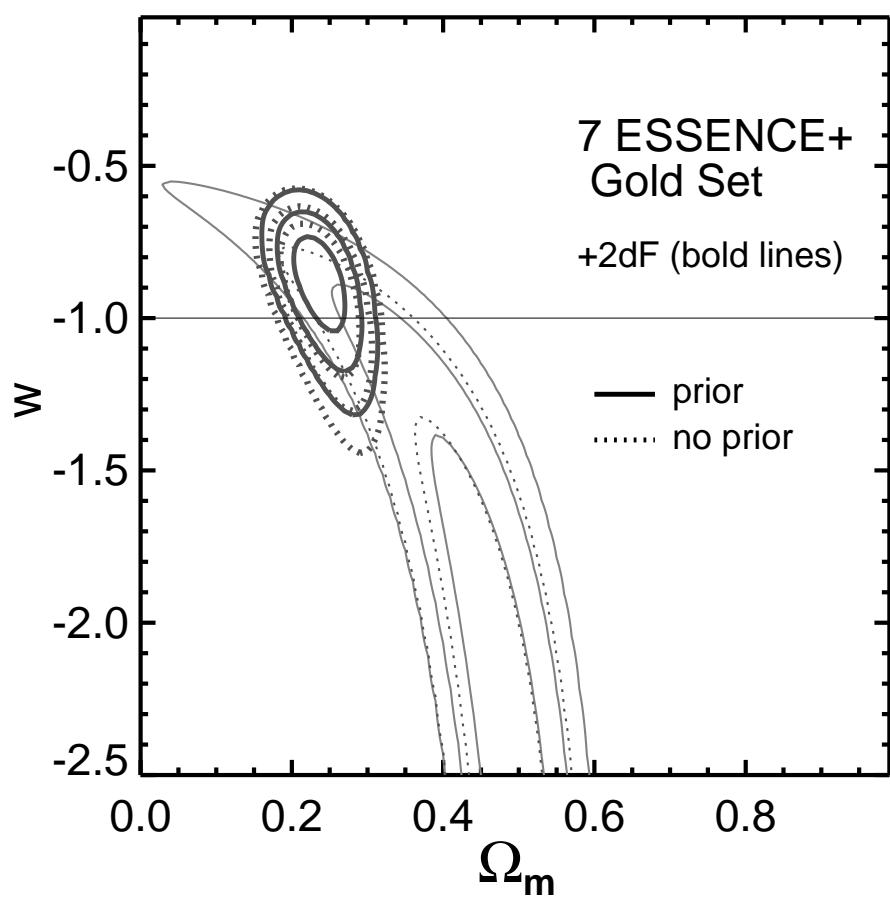
Krisciunas *et al.* Fig. 13



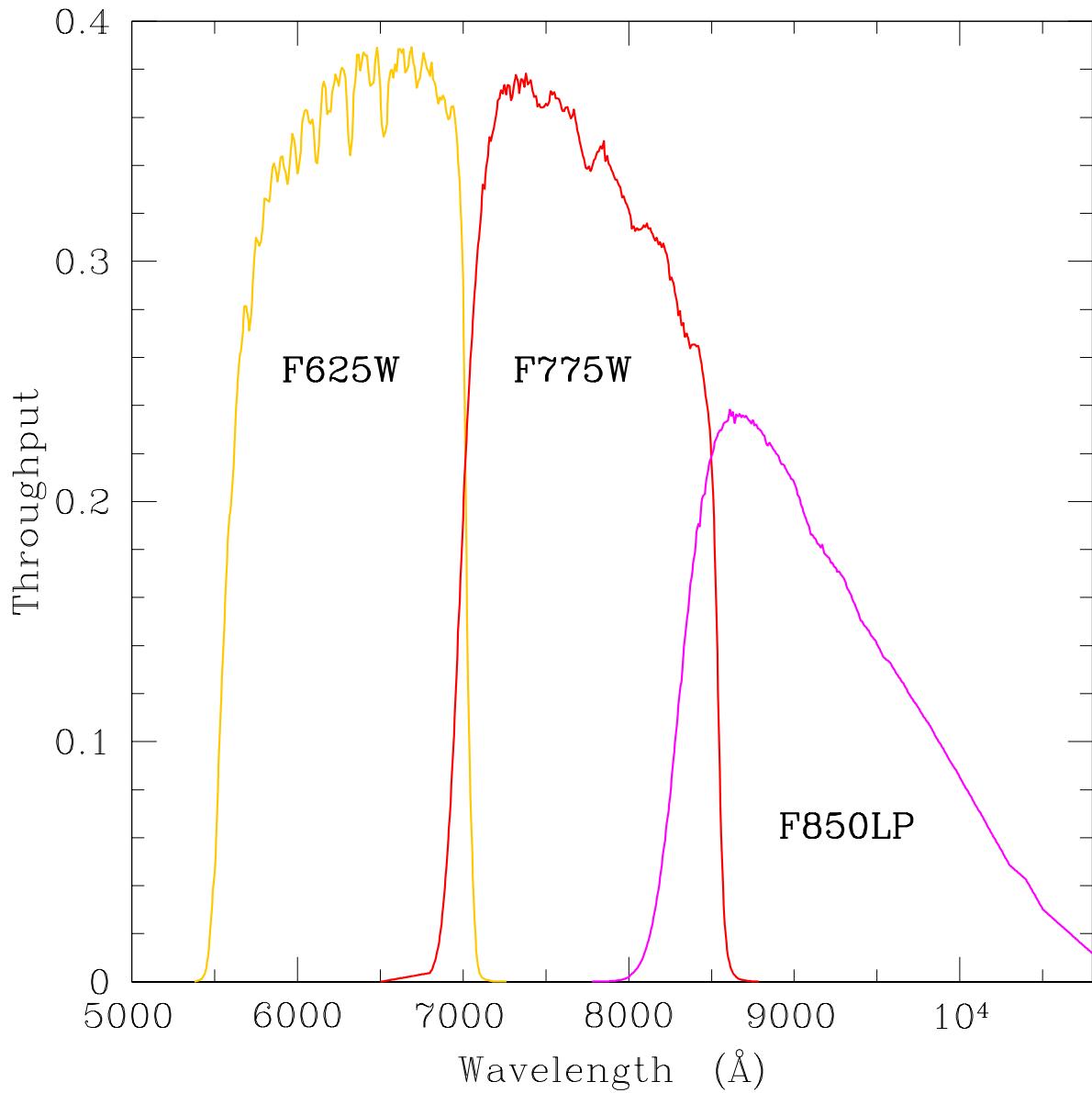
Krisciunas *et al.* Fig. 14



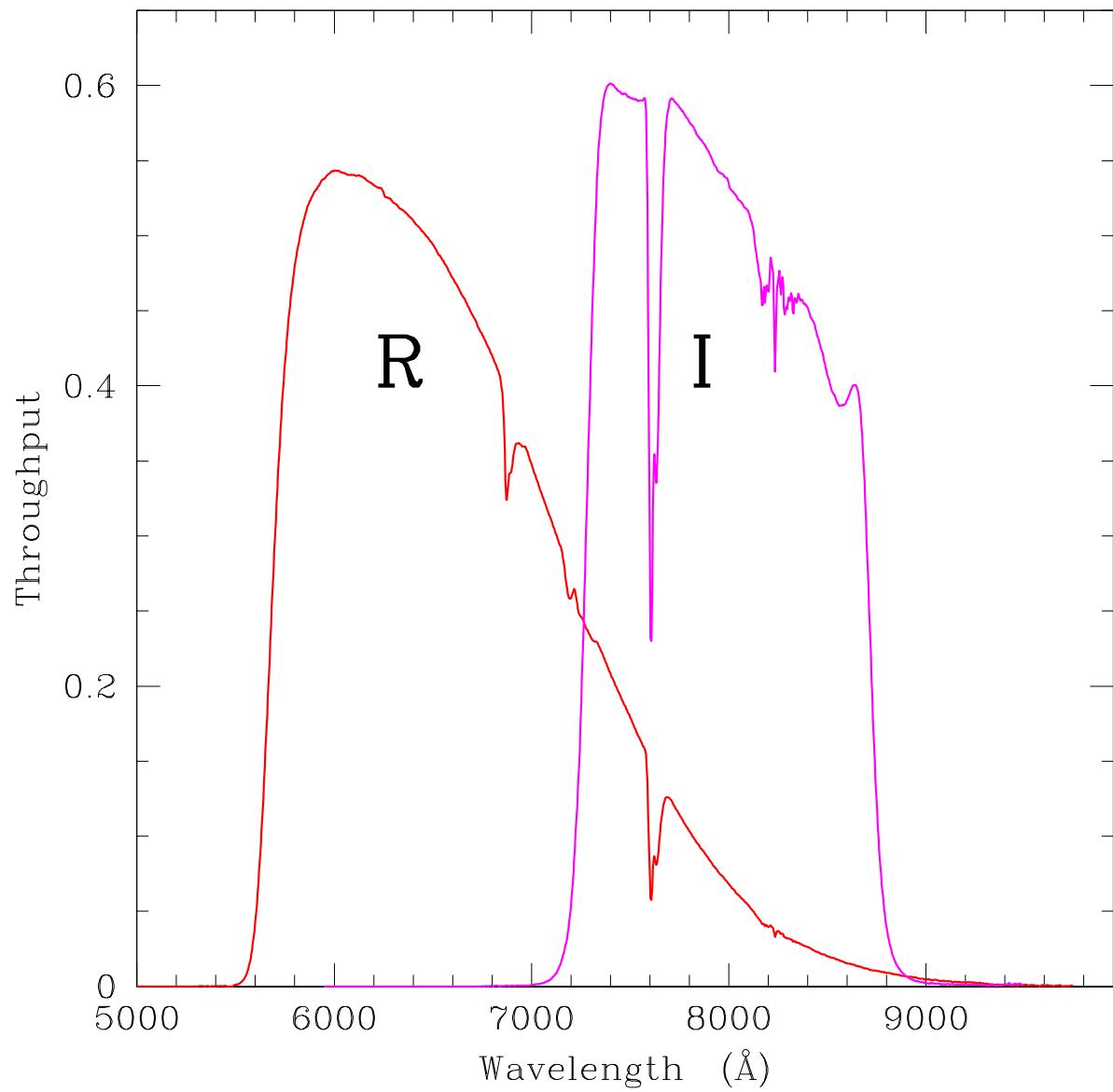
Krisciunas *et al.* Fig. 15



Krisciunas *et al.* Fig. 16



Krisciunas *et al.* Fig. 17



Krisciunas *et al.* Fig. 18